

# Identification of active galaxies behind the Coma cluster of galaxies<sup>★</sup>

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**Abstract.** We describe the results of an efficient survey to identify bright  $B \lesssim 18$  active galactic nuclei (AGN) within  $\approx 4^\circ$  ( $4.7 h^{-1}$  Mpc projected radius) of the centre of the Coma cluster of galaxies. The candidates have been carefully selected either as unresolved UV-excess objects with no detected proper motion, or from identifications on objective-prism plates including a few ROSAT sources. A global success rate larger than 30% is achieved. Low-resolution spectra are presented for the extragalactic objects identified in the course of the survey: the sample contains 16 emission-line objects, a BL Lac candidate and an elliptical galaxy. The AGNs of this list are unique targets to study the association of the Lyman- $\alpha$  forest with large-scale structures and galactic haloes at low redshift, to search for warm gas within the Coma cluster and nearby filaments, and for optical monitoring to detect the microlensing variability produced by the baryonic dark matter in the cluster.

**Key words:** quasars: general — quasars: absorption lines — galaxies: Seyfert — galaxies: starburst — techniques: spectroscopic — surveys

## 1. Introduction

Bright quasars (QSOs) can be used as background sources to probe the physical properties of foreground objects.

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<sup>★</sup> Based on observations carried out at the 1.93 m telescope of the Observatoire de Haute-Provence (CNRS), France.

When the foreground is a cluster of galaxies, the quasars behind the cluster can be used to map not only the warm gas content of the intra-cluster medium (via absorption lines), but also its dark matter distribution (via microlensing).

The discovery of a diffuse extreme UV-excess in nearby clusters by the *Extreme Ultraviolet Explorer* (EUVE) satellite (Lieu et al. 1996a,b) was initially interpreted as evidence for a warm gas component at about 500 000 K in the intra-cluster medium. This would constitute a major mass component and raises enormous problems (Fabian 1996). A non-thermal explanation, based on inverse Compton scattering, seems more plausible and consistent with the hard X-ray tail observed by *BeppoSAX* (Ensslin & Biermann 1998; Sarazin & Lieu 1998). However the distorted shape of the EUV isophotes may be more consistent with a shock heated medium, so the situation remains very confused. Although attempts have been made to detect the warm gas through UV emission lines (Dixon et al. 1996), absorption lines are far more sensitive to the presence of this component. Hence, *Hubble Space Telescope* (HST) spectroscopy of background QSOs should reveal absorption lines at the redshift of the cluster if the thermal interpretation is correct.

In addition, background QSOs can be used as probes of intervening gaseous clouds and of their evolution along the line of sight. Adjacent lines of sight provided by close pairs of QSOs can also constrain the size of the absorbers, their spatial distribution and their connection with galaxies. In this context, targeting quasars behind low-redshift clusters is a very efficient way to reveal the connection of the intergalactic medium on any scale with these

concentrations of galaxies. Coma (Abell 1656, at a redshift of  $\langle z_{\text{Coma}} \rangle = 0.0229$ ) is ideal in this respect since it is one of the best studied areas of the sky as far as redshift coverage and galaxy properties are concerned (Biviano 1998).

Quasar pairs with projected separations of no more than a few arcminutes yield interesting constraints on the size, physical structure and kinematics of galactic haloes, clusters and filaments (Smette et al. 1995; Petitjean et al. 1998; D’Odorico et al. 1998). The ideal project however would be to observe a large number of QSOs in a small solid angle on the sky to probe all scales at the same time. Since absorption is extremely sensitive to the total amount of warm gas, this would probe even small column densities inside the cluster and also reveal, in the outskirts of the cluster, the connection with the large-scale filamentary network (Cen et al. 1994; Petitjean et al. 1995; Hernquist et al. 1996; Miralda-Escudé et al. 1996).

Besides the mapping of the warm gas in the cluster and in the neighbouring environment, a third goal of a QSO survey behind clusters is the mapping of their baryonic dark matter content. This can be achieved by monitoring the background QSOs searching for the expected micro-lensing signature in their light curves. The micro-lensing optical depth produced by a cluster is very large, a few  $10^{-3}$ , and for lens masses ranging from  $10^{-5}$  to  $10^{-3}M_{\odot}$  the events will have typical time scales of days (Walker & Ireland 1995; Tadros et al. 1998; Wu & Xue 1998). A program to monitor a few tens of QSOs behind Coma for instance, for 5 months with a daily sampling rate, should reveal dozens of micro-lensing events and produce a 2-D mapping of the baryonic dark matter within the cluster.

There are however very few quasars known to date behind low-redshift clusters and clearly additional background quasars, and more generally galaxies with an active galactic nucleus, are badly needed for all these studies. These sources have to be bright enough to be accessible for medium-resolution spectroscopy with the HST and for optical monitoring with small/medium-size ground-based telescopes.

Until now optical searches for quasars have only been performed systematically behind Virgo (He et al. 1984; Impey & He 1986), where 29 QSOs were found out of 82 candidates observed at the Palomar 5-meter telescope. There are a few X-ray selected QSOs found in ROSAT images whose optical counterparts were identified in the *Digitized Sky Survey* around several clusters (Knezek & Bregman 1998). Although interesting, this technique only provides a handful of QSOs in the vicinity of clusters. There are also unpublished surveys around Abell 2029 and Abell 2255 (Horowitz et al. 1994). We report here on a survey (Ledoux et al. 1998) to complete the extant catalogues of QSOs/AGNs within a projected radius of  $\approx 4^{\circ}$  from the centre of the Coma cluster,

down to a limiting magnitude of about  $B < 18$ . We describe the selection methods in Sect. 2 and present the observations in Sect. 3. In Sects. 4 and 5, we discuss the characteristics of the extragalactic objects identified in this way, and conclude in Sect. 6 on the global efficiency of the selection.

Throughout this paper, we assume a standard Einstein-De Sitter cosmological model and a Hubble constant  $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

## 2. Target selection

AGN candidates were selected with the criterion  $B \lesssim 18$  from 3 different sources: [1] the OMHR survey of UV-excess objects at the North Galactic pole (Moreau & Reboul 1995), [2] the Hamburg Quasar Survey (HQS) by Hagen et al. (1995) and [3] the *ROSAT All-Sky Survey* (RASS) using identifications on HQS plates (Bade et al. 1998).

We also selected the object US 370 (Usher 1981), which was classified as a QSO candidate with  $V = 18$  and  $U - V = -0.55$  by Berger et al. (1991). The proper motion of this target could not be measured due to the proximity of a bright star as mentioned by Moreau & Reboul (1995). A slit-less observation was performed by Weedman (1985) and we provide here a spectrum of US 370 to confirm its nature and measure its redshift properly.

### 2.1. OMHR candidates

Berger et al. (1991) performed a photometric ( $U, B, V$ ) and astrometric analysis of 1221 UV-excess unresolved objects in a  $40.5^{\square}$  field almost centred on the selected area SA 57:

$$12^{\text{h}}52^{\text{m}}08.6^{\text{s}} < \alpha_{\text{J2000}} < 13^{\text{h}}21^{\text{m}}39.2^{\text{s}} \\ 26^{\circ}01'12'' < \delta_{\text{J2000}} < 32^{\circ}14'39''.$$

They used the MAMA (*Machine Automatique à Mesurer pour l’Astronomie*) machine to scan three plates taken in 1962 at the Palomar 48-inch Schmidt. This field was fully reduced in  $U, B$  and  $V$  by Moreau & Reboul (1995) with a visual limiting magnitude 20.3. Amongst the  $\sim 130\,000$  detected objects, the authors extracted a list of 1759 quasar candidates selected mainly as UV-excess unresolved objects. The colour criterion was  $U - V < 0.1$  with a typical accuracy of 0.1 for the  $U$  and  $V$  magnitudes. After rejection of the objects with previous spectroscopic identifications, a total of 1681 quasar candidates remained.

A recent comparison of the MAMA catalogues extracted from the Schmidt plates taken in 1962 with those obtained in the same way from OCA (Observatoire de la Côte d’Azur) Schmidt plates taken in 1991 and 1994 provides measures of proper motions of the above OMHR quasar candidates, among many other objects (Moreau & Reboul, in preparation). The limited field of view of the

OCA telescope constrained this proper motion survey to a  $26.7''$  field within:

$$12^{\text{h}}56^{\text{m}}42.1^{\text{s}} < \alpha_{\text{J2000}} < 13^{\text{h}}20^{\text{m}}41.4^{\text{s}} \\ 26^{\circ}50'42'' < \delta_{\text{J2000}} < 31^{\circ}57'51''.$$

The  $1\sigma$  sensitivity of the measurement of proper motions is 5 to 8  $\text{mas yr}^{-1}$  for the magnitude range  $16 < V \lesssim 18$  used in the present work.

The combined requirements of [1] unresolved morphology, [2] UV-excess and [3] absence of proper motion larger than about 10  $\text{mas yr}^{-1}$  (i.e. a transverse velocity less than 25  $\text{km s}^{-1}$  at 500 pc) give a priori high-quality quasar candidates. However, the constraint imposed by medium resolution spectroscopy with the HST restricted our choice to objects brighter than magnitude 18. Quasars this bright are drastically scarcer (and consequently samples are far more polluted by stars) than for visual magnitudes fainter than  $V = 18$  where the best and most numerous quasar candidates are expected to lie.

The sample of OMHR objects selected for observation finally contained the 18 brightest and bluest targets ( $B \lesssim 18$ ,  $U - V \leq 0$ ) with no proper motion detected at the 0.68 confidence level.

## 2.2. HQS and RASS candidates

Quasar candidates were selected from the HQS digitized objective-prism plates in two steps (Hagen et al. 1995, 1999). In the first step, prism spectra with a blue continuum are selected automatically from the digitized data which are available only at low resolution. The selected spectra are rescanned with the Hamburg PDS machine to yield high-resolution spectra having a dispersion of 139  $\text{nm mm}^{-1}$  at  $\text{H}\gamma$ . These spectra are interactively classified on a vector graphics screen. Non-blue spectra with stellar absorption features are discarded while the others are kept in an archive.

Blue stellar-like spectra are classified as “QSO/HS”, meaning quasar or hot star candidate. Spectra with apparent emission lines or unusual continua are classified as “QSO” or “Narrow Emission”. If the spectra are blue and featureless they enter the category “Unid” for “unidentified”. The number of entries in these classes increases towards the “unidentified” sources, while the efficiency to find quasars in these classes decreases from “QSO” towards the other classes.

We selected from the archive all the objects with  $B \leq 18$  contained in the above mentioned classes within about  $4^\circ$  from the centre of the Coma cluster. Hagen et al. (1999) estimate that for  $z > 0.1$  and  $B \leq 17.5$  less than 10% of the quasars are missed. Objects already known from the literature were discarded by cross-correlation with the NED and SIMBAD databases. We checked for overlaps with the OMHR survey, and discarded as well the objects with a significant proper motion (confidence

level  $> 0.68$ ) or colour  $U - V > 0$ . The final HQS list used to obtain follow-up spectroscopy contains 31 objects.

The HQS list was augmented by two AGN candidates from the RASS. X-ray sources from this survey are identified on the Schmidt plates of the HQS. The confirmation rate of the AGN candidates is very high, about  $> 90\%$  (Bade et al. 1992), but most of them are optically fainter than our selection limit or were already known. The first RASS source RXJ1252.6+3002 will be contained in the RASS Faint Source Catalog (Voges et al., in preparation) while the second one RXJ1303.7+2633 is published in the RASS Bright Source Catalogue (Voges et al. 1997). Also, HS 1312+2735 is listed as X-ray source in the latter catalogue (namely RXJ1314.3+2719).

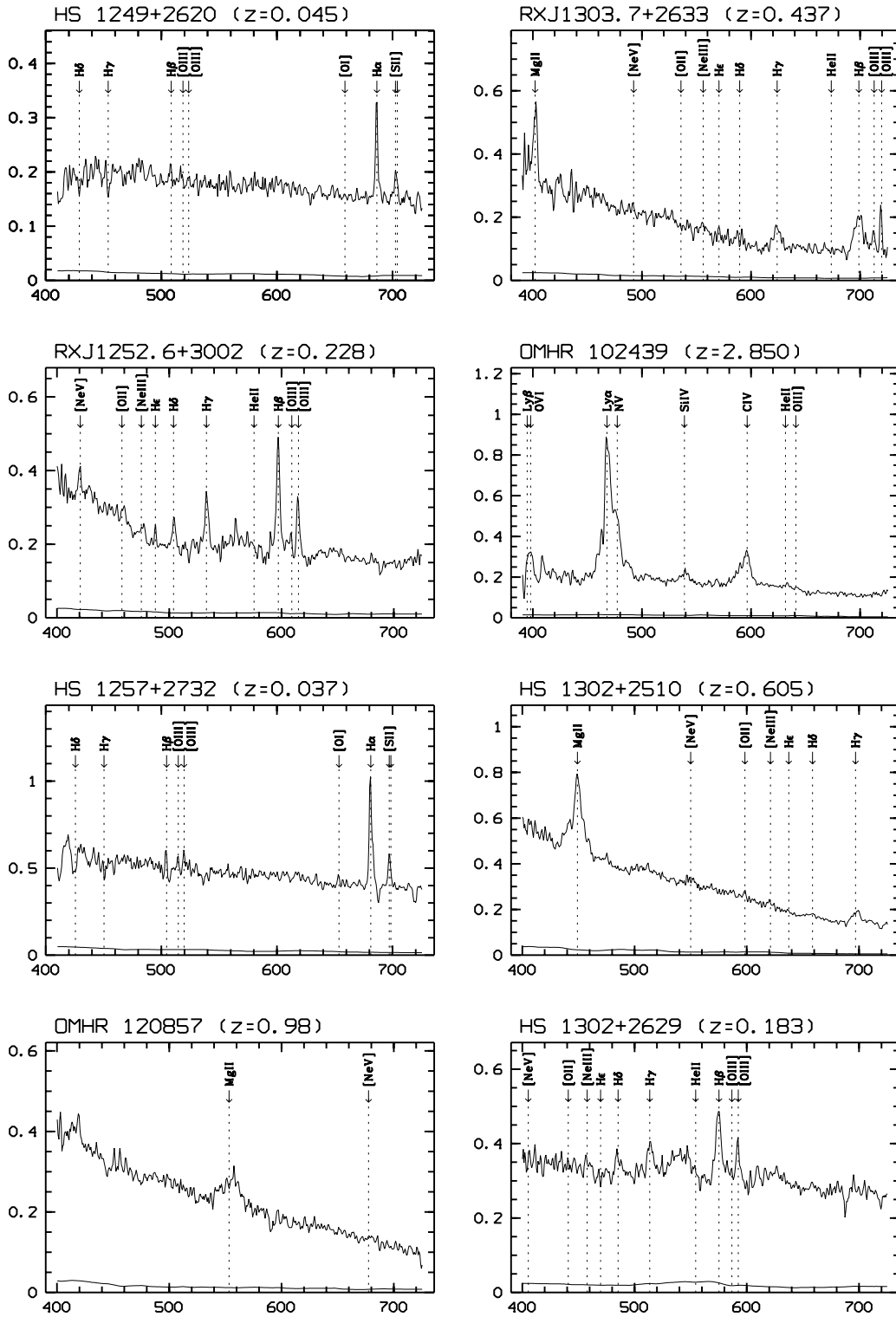
## 3. Optical spectroscopy

The spectroscopic observations aimed at confirming the nature of the selected candidates were carried out at the F/15 Cassegrain focus of the 1.93 m telescope at the Observatoire de Haute-Provence (CNRS, France) during the nights of April 8-15, 1997 and April 20-21, 27-29, 1998. The long-slit spectrograph CARELEC (Lemaitre et al. 1990) was operated with a 150 lines  $\text{mm}^{-1}$  grating leading to a dispersion of 26  $\text{nm mm}^{-1}$  in the wavelength range  $\lambda\lambda 380.0 - 725.0$  nm. The detector was a  $512 \times 512$  Tektronix CCD with 27  $\mu\text{m}$  pixels and a rms read-out noise of 9.2 electrons. Throughout the observing runs the atmospheric seeing usually ranged from  $1''.5$  to  $2''.0$ . Most of the observations were made with a  $2''.5$  slit width, which corresponds to a projected slit width on the detector of 62  $\mu\text{m}$  or 2.3 pixels. With this set-up, the overall instrumental resolution FWHM was 1.4 nm as measured on the He arc lines used for the wavelength calibration. An exposure time of 3600 s was enough to yield a signal-to-noise ratio larger than 10 in the spectra of the faintest candidates ( $B \approx 18$ ).

The sub-dwarf O star HZ 44 (Oke 1990) was used as a spectro-photometric standard during both observing runs. Depending on the atmospheric stability of the night, the absolute flux calibration is uncertain at the 20 to 50% level. The data were reduced and analyzed independently by two of us, with the LONG-SLIT package implemented within MIDAS (CL), and with IRAF (DVG). After flat-field correction and wavelength calibration, the object signal was extracted using the Horne algorithm (Horne 1986). It ensures an optimal signal-to-noise ratio for weak spectra while it possibly slightly underestimates the emission-line strength in extended sources.

## 4. Sample characteristics

The coordinates, apparent magnitudes and redshifts of the extragalactic objects are listed in Table 1.



**Fig. 1.** Optical spectra of the extragalactic objects at a resolution FWHM of 1.4 nm: the calibrated flux  $F_\lambda$  in units of  $10^{-17} \text{ W m}^{-2} \text{ nm}^{-1}$  is plotted against heliocentric wavelength  $\lambda$  in nm. The curve below each spectrum is the associated  $1\sigma$  standard deviation of the noise in the same flux units

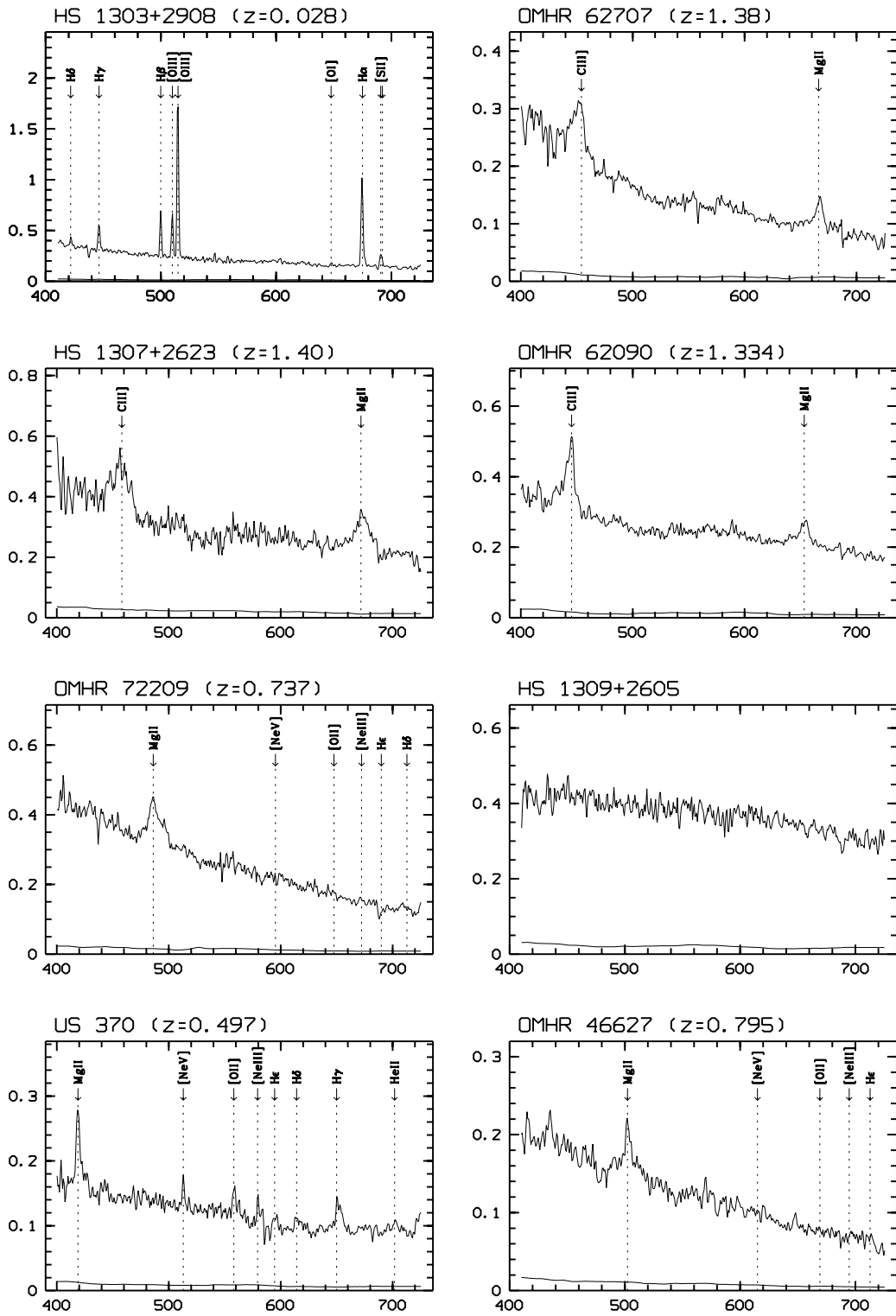


Fig. 1. continued

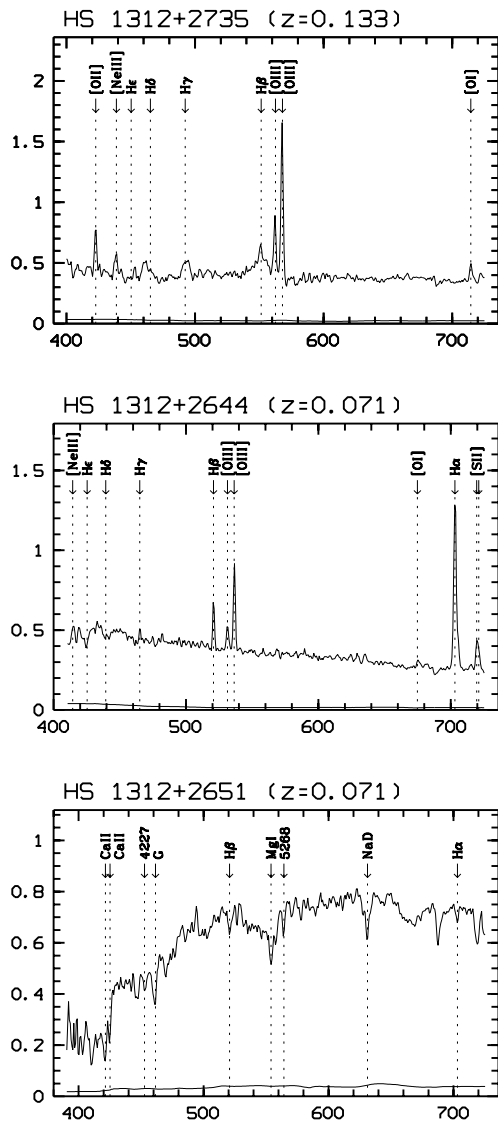


Fig. 1. continued

The projected impact parameters (Col. 6) range from 0.8 to about  $4.7 h^{-1}$  Mpc, and therefore probe the cluster from the inner Abell radius all the way out to the low-density environment and large-scale filaments. The wavelength and flux calibrated 1-D spectra are displayed in Fig. 1. The mean signal-to-noise ratio has been computed as a function of wavelength from the photon statistics after subtraction of the spectral lines in the normalized spectrum. Besides the stellar objects (which are dealt with in a separate paper), we discovered 16 new emission-line objects, a BL Lac candidate and an elliptical galaxy. We confirmed the redshift of the QSO US 370. The emission-line objects consist of 9 quasars, 3 Seyfert 1 and 4 HII-region galaxies (see below). We included HS 1309+2605 as a BL Lac candidate because none of the usual stellar absorption features could be detected in the slit spectrum (see Fig. 1). However there has been no radio detection (Kim et al. 1994; Kim 1994) and the signal-to-noise ratio in our optical spectrum is insufficient

for a final decision. The normal galaxy HS 1312+2651 was originally classified as a ‘‘Narrow Emission’’ object according to the objective-prism spectrum, but the emission feature could not be confirmed.

Twelve emission-line objects show broad permitted emission lines and should be quasars or Seyfert 1 galaxies. We adopt the definition by Véron-Cetty & Véron (1998) of a quasar as a star-like object, or an object with a star-like nucleus, brighter than the absolute  $B$ -band magnitude of  $-21.5 + 5 \log h$ . We have made a distinction between QSOs and Seyfert 1 galaxies on the basis of this definition, and calculated the absolute  $B$ -band magnitude  $M_B$  of all the objects in Table 1 (Col. 10). Note that no  $K$ -correction was applied since it is usually of the same order as the uncertainties in the photometry and/or it does not affect the classification.

Among the Seyfert galaxies, we find several sub-types. RXJ1252.6+3002 belongs to the class of narrow-line Seyfert 1 galaxies (Boller et al. 1996). The intrinsic FWHM of  $H\beta$ , around  $1150 \text{ km s}^{-1}$ , is typical for such objects, and it is only slightly larger than that of  $[\text{OIII}]\lambda 500.7$ . Moreover, the flux ratio of  $[\text{OIII}]\lambda 500.7$  to  $H\beta$  is less than unity which rules out the possibility that it is a Seyfert 2 (Shuder & Osterbrock 1981). Additional support for this classification comes from the detection of strong blends of FeII lines in the regions around 457.0 nm and 519.0 nm (Koski 1978). HS 1302+2629 also shows relatively narrow lines, but given its FWHM at  $H\beta$  of  $\approx 2350 \text{ km s}^{-1}$ , it has to be classified as a classical Seyfert 1 galaxy. Although HS 1312+2735 is obviously a Seyfert 1 galaxy, a Seyfert 1.5 type (Netzer 1990) is still possible given the shape of the  $H\beta$  line.

The intrinsically faintest objects HS 1249+2620, HS 1257+2732, HS 1303+2908 and HS 1312+2644 only show narrow permitted emission lines. In Sect. 5 below, we show that these narrow emission-line galaxies (NELGs) are indeed HII region (i.e. starburst) galaxies.

## 5. Redshifts and line parameters

Due to the large wavelength coverage of the spectra, more than one emission line is detected in most of the cases of objects classified as QSOs, Seyfert 1 galaxies or NELGs, providing secure identifications. When possible three prominent emission lines were selected, preferentially unresolved features far from the border of the spectra where the wavelength calibration is possibly less reliable. A Gaussian profile was fitted to those lines (see Table 1, Col. 8) so as to determine their centroid, and the AGN’s redshift was computed by averaging the redshift of the individual lines. In Table 1 (Col. 9), redshift values are vacuum and given in the heliocentric rest frame. The formal error on them is typically of the order of 0.1 nm or  $\Delta z \sim 10^{-3}$ , except for OMHR 120857 and HS 1307+2623 where one (or two) broad and somewhat asymmetric line(s) lead(s) to a larger centering error (up to  $\Delta z = 0.01$ ).

**Table 1.** The sample of confirmed extragalactic sources

Nr	Object	$\alpha$ (J2000) (h m s)	$\delta$ (J2000) ( $^{\circ}$ ' ")	$\theta_{\text{Coma}}$ ( $^{\circ}$ )	$\rho_{\text{Coma}}^{\text{a}}$ ( $h^{-1}\text{Mpc}$ )	$B^{\text{b}}$	Lines	$z_{\text{helio}}$	$M_B^{\text{c}}$ $-5 \log h$	Type
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
1	HS 1249+2620	12 51 54.4	+26 03 59	2.58	3.0	17.5	H $\beta$ ,H $\alpha$	0.045	-18.2	HII
2	RXJ1252.6+3002	12 52 37.8	+30 02 13	2.58	3.0	18.2	H $\gamma$ ,H $\beta$ ,[OIII]	0.228	-21.1	NLSy1
3	HS 1257+2732	13 00 14.4	+27 16 49	0.70	0.8	16.0	[OIII],H $\alpha$	0.037	-19.2	HII
4	OMHR 120857	13 02 03.1	+26 58 09	1.13	1.3	17.1	MgII	0.98:	-25.6	QSO
5	RXJ1303.7+2633	13 03 46.0	+26 33 11	1.68	1.9	17.8	MgII,H $\gamma$ ,H $\beta$	0.437	-23.0	QSO
6	OMHR 102439	13 04 12.0	+29 53 49	2.16	2.5	18.1	Ly $\alpha$ ,SiIV,CIV	2.850	-27.2	QSO
7	HS 1302+2510	13 04 51.3	+24 54 47	3.27	3.8	17.4	MgII,[OII]	0.605	-24.1	QSO
8	HS 1302+2629	13 05 16.7	+26 13 03	2.15	2.5	17.6	H $\gamma$ ,H $\beta$ ,[OIII]	0.183	-21.2	Sy1
9	HS 1303+2908	13 05 45.1	+28 52 12	1.60	1.8	16.8	H $\beta$ ,[OIII]	0.028	-17.8	HII
10	HS 1307+2623	13 09 28.8	+26 07 48	2.85	3.3	17.7	CIII],MgII	1.40	-25.8	QSO
11	OMHR 72209	13 09 44.9	+28 27 35	2.26	2.6	17.9	MgII	0.737:	-24.1	QSO
12	US 370 <sup>d</sup>	13 10 28.1	+28 08 38	2.38	2.7	17.9	MgII,[NeV],[OII]	0.497	-23.2	QSO
13	OMHR 62707	13 10 59.0	+30 01 07	3.20	3.7	18.1	CIII],MgII	1.38	-25.4	QSO
14	OMHR 62090	13 11 28.4	+31 05 51	4.04	4.7	18.0	CIII],MgII	1.334	-25.4	QSO
15	HS 1309+2605	13 12 19.2	+25 49 58	3.53	4.1	17.8	...	...	...	BL Lac:
16	OMHR 46627	13 13 36.2	+29 58 15	3.64	4.2	18.1	MgII	0.795:	-24.1	QSO
17	HS 1312+2735	13 14 23.5	+27 19 20	3.31	3.8	17.2	[OIII]	0.133	-20.9	Sy1
18	HS 1312+2644	13 14 31.9	+26 28 31	3.62	4.2	17.0	H $\beta$ ,[OIII]	0.071	-19.7	HII
19	HS 1312+2651	13 15 02.1	+26 35 43	3.67	4.2	17.4	H $\beta$ ,MgI,H $\alpha$	0.071	-19.3	E

<sup>a</sup> Adopting  $\langle z_{\text{Coma}} \rangle = 6853 \text{ km s}^{-1}$ .

<sup>b</sup> The magnitudes of the HQS and RASS sources were determined on the objective-prism plates and may have errors up to 0.5 mag; the OMHR sources have typical errors in  $B$  of 0.1 mag.

<sup>c</sup> No K-correction is applied.

<sup>d</sup> Usher (1981).

Table 2 displays the measurements of the line parameters for the entire sample of emission-line objects. The numbers of Col. 1 are the same as in Col. 1 of Table 1. The equivalent width  $W_{\lambda_r}$  of the detected lines (Col. 3) and their full width at half-maximum FWHM (Col. 4) are both corrected from redshift (rest frame values); FWHM is in addition corrected for the instrumental response. No FWHM value is given for unresolved lines.

The classification of the NELGs can be made on the basis of diagnostic diagrams which discriminate between different physical processes (Baldwin et al. 1981): either photoionization by a non-thermal power-law continuum in a narrow-line AGN (i.e. Seyfert 2 galaxies and LINERS), or photoionization due to UV photons emitted by hot OB stars (i.e. HII region galaxies). We have computed *flux ratios* for the most important diagnostic lines (Veilleux & Osterbrock 1987). Criteria involving ratios between adjacent lines were used to get rid as much as possible of extinction problems. From the [OIII]/H $\beta$  - [SII]/H $\alpha$  diagram, all 4 NELGs turn out to be unambiguously objects with a HII-like spectrum. The diagnostics based on [OIII]/H $\beta$  - [NII]/H $\alpha$  and [OIII]/H $\beta$  - [OI]/H $\alpha$  diagrams are a bit less certain given the errors in the measured fluxes, but give results which are consistent with this classification.

## 6. Discussion

The selection of HQS candidates led to an efficiency of  $29\% \pm 10\%$  taking into account the probable BL Lac object HS 1309+2605. Similarly, the  $U - V \leq 0$  OMHR candidates with undetected proper motion gave a success rate for identified QSOs of  $33\% \pm 14\%$ .

The OMHR selection actually gives a success rate of  $54\% \pm 14\%$  if the 8 previously known QSOs falling under the same selection criteria are included, which is a good score for this range of (bright) optical magnitudes. Similarly, deep pointed ROSAT surveys towards low galactic neutral hydrogen column density regions (Shanks et al. 1991; Hasinger et al. 1993, 1994) give a selection of AGNs in X-rays that is very successful with typical efficiencies of up to  $\sim 60\%$  (Zickgraf et al. 1997). We also note that the fact that all OMHR targets are either stars or *bona fide* QSOs confirms the quality of the diffuse-object rejection by Moreau & Reboul (1995).

Out of a total of 51 selected candidates, 16 new active galaxies (AGNs and HII region galaxies) have been identified, and a probable BL Lac object discovered. This leads to a  $33\% \pm 8\%$  global selection efficiency. Compared to the catalogue of Véron-Cetty & Véron (1998), the number of confirmed QSOs/AGNs brighter than 18<sup>th</sup> magnitude<sup>1</sup> has increased by 22% within the inner 2 $^{\circ}$ ,

<sup>1</sup> Note however that some sources are ambiguous, like BG 57 27, BG 57 49 and BG 57 42.

**Table 2.** Emission-line measurements

Nr	$\lambda_{\text{helio}}$ (nm)	$W_{\lambda_r}$ (nm)	$FWHM$ (km s <sup>-1</sup> )	$S/N$	Transition	Nr	$\lambda_{\text{helio}}$ (nm)	$W_{\lambda_r}$ (nm)	$FWHM$ (km s <sup>-1</sup> )	$S/N$	Transition		
(1)	(2)	(3)	(4)	(5)	(6)	(1)	(2)	(3)	(4)	(5)	(6)		
1	508.0	0.4	.....	16	H $\beta$	9	499.8	3.0	.....	21	H $\beta$		
1	523.5	0.2	987	16	[OIII]	500.7	9	509.9	3.4	.....	20	[OIII]	495.9
1	686.2	3.1	550	20	H $\alpha$	9	514.7	11.7	.....	18	[OIII]	500.7	
1	702.9	1.1	1023	20	[SII]	672.5	9	674.5	11.1	.....	25	H $\alpha$	
2	420.3	0.4	1308	14	[NeV]	342.6	9	691.0	1.8	733	24	[SII]	672.5
2	504.2	1.0	1264	16	H $\delta$	10	456.2	3.2	7306*	15	CIII]	1909	
2	533.4	2.4	1549	15	H $\gamma$	10	673.0	2.1	2101	18	MgII	279.8	
2	597.0	6.8	1146	15	H $\beta$	11	486.1	2.9	5948	34	MgII	279.8	
2	614.7	2.7	1090	20	[OIII]	500.7	12	418.9	3.5	1628	13	MgII	279.8
3	504.0	0.6	746	13	H $\beta$	12	513.0	0.6	.....	20	[NeV]	342.6	
3	519.5	0.4	842	13	[OIII]	500.7	12	558.5	0.8	858	14	[OII]	372.7
3	680.9	4.1	1040	31	H $\alpha$	12	579.8	0.5	.....	13	[NeIII]	386.9	
3	697.2	1.1	829	25	[SII]	672.5	12	615.1	0.8	2041	12	H $\delta$	
4	553.4	3.1	6368	15	MgII	279.8	12	651.8	2.1	2018	19	H $\gamma$	
5	402.2	2.3	1912	14	MgII	279.8	13	452.5	2.3	4206	27	CIII]	190.9
5	623.8	3.8	2561	17	H $\gamma$	13	666.8	1.0	1147	20	MgII	279.8	
5	698.9	8.0	2728	12	H $\beta$	14	445.5	3.0	2641	25	CIII]	190.9	
5	712.2	1.5	771	12	[OIII]	495.9	14	653.3	0.9	1586	30	MgII	279.8
5	718.9	2.4	431	12	[OIII]	500.7	16	502.5	2.0	1774	19	MgII	279.8
6	397.5	1.0	871	12	Ly $\beta$	17	422.7	1.6	.....	13	[OII]	372.7	
6	468.2	10.5	2268*	16	Ly $\alpha$	17	438.5	1.1	1344	14	[NeIII]	386.9	
6	539.2	0.4	1179	22	SIV	140.0	17	461.5	1.5	3277	16	H $\delta$	
6	595.7	2.3	974	19	CIV	154.9	17	493.2	2.1	2878	14	H $\gamma$	
7	449.2	4.7	4028*	35	MgII	279.8	17	551.3	3.6	3485	19	H $\beta$	
7	598.0	0.1	662	21	[OII]	372.7	17	562.1	2.6	.....	18	[OIII]	495.9
7	697.8	1.1	2020	16	H $\gamma$	17	567.7	5.8	.....	17	[OIII]	500.7	
8	485.0	0.7	2619	19	H $\delta$	17	714.6	0.7	667	13	[OI]	630.0	
8	514.1	1.3	2521	17	H $\gamma$	18	465.2	0.3	.....	15	H $\gamma$		
8	574.8	3.0	2347	18	H $\beta$	18	520.8	1.3	.....	28	H $\beta$		
8	591.9	0.9	1137	22	[OIII]	500.7	18	531.4	0.8	.....	29	[OIII]	495.9
9	397.7	0.6	.....	19	[NeII]	386.9	18	536.5	2.3	.....	31	[OIII]	500.7
9	421.9	0.3	.....	19	H $\delta$	18	703.2	9.7	697	27	H $\alpha$		
9	446.3	1.6	.....	22	H $\gamma$	18	720.2	1.9	852	33	[SII]	672.5	

\* Gaussian fit inappropriate/impossible for this line.

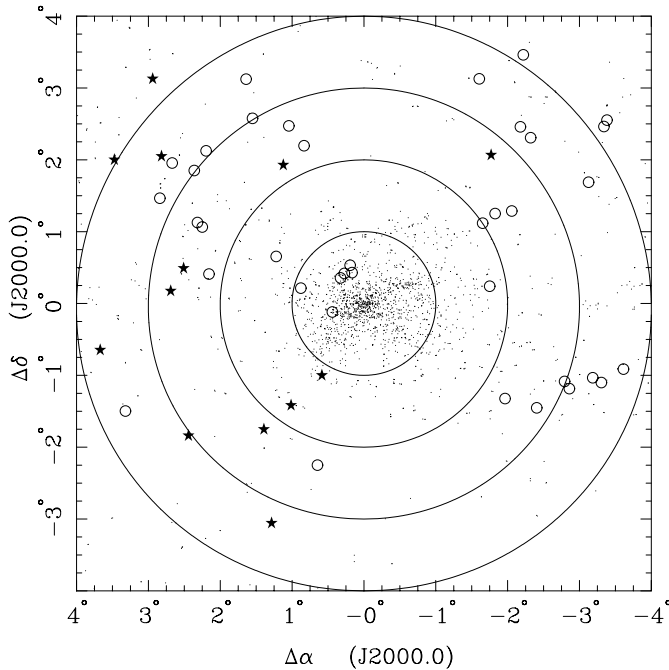
and by 36% in the annulus between 2° and 4° (see Fig. 2).

No less than 50 QSOs/AGNs with  $B$  or  $V \lesssim 18$  are now available behind Coma within a radius of  $\approx 4^\circ$  of its centre, and they constitute the largest sample of QSOs/AGNs available behind a cluster, with a more uniform coverage from the inner Abell radius to the outskirts related to the large-scale filaments. Figure 2 gives their angular distribution and the positions of galaxies with known redshifts from a variety of surveys (Karachentsev & Kopylov 1990; Huchra et al. 1995; Biviano et al. 1995; Doi et al. 1995; Donas et al. 1995; Colless & Dunn 1996). Note that there are far many more galaxies known in this field (Godwin & Peach 1977; Slezak et al. 1988), whose redshifts are likely to be measured in the near future by the *Sloan Digital Sky Survey*. Moreover, there are  $\sim 15$  QSO/AGN pairs with separations smaller than 30'. The sample thus provides ideal targets for observation with HST and for optical

monitoring with small/medium-size telescopes to shed some light on the nature of the Lyman- $\alpha$  clouds at low redshift, to study the spatial distribution of the warm gas in and around the Coma cluster of galaxies, and its baryonic dark matter content.

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**Fig. 2.** Projected distribution of the new AGNs brighter than  $B \approx 18$  behind Coma (filled stars) together with the AGNs listed in the compilation by Véron-Cetty & Véron (1998) brighter than  $V \approx 18$  in the same redshift range (empty circles). The dots indicate galaxies with measured redshifts within that solid angle (see text)

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