

Photometric and morphological analysis of UCM galaxies.

II. Statistical results and interpretation*

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Abstract. — We present the results of Thuan & Gunn r CCD imaging observations of the Universidad Complutense de Madrid (UCM) sample of emission-line galaxies (ELGs) selected by the presence of $H\alpha$ emission in low resolution objective-prism spectra. In this work, a total of 212 objects from the UCM survey have been photometrically and morphologically characterized. In a previous paper, the observations, basic reductions and morphological classification for the UCM objects were presented. Here, we analyse the statistical properties of the sample and their interpretations. The main results obtained are as follows: (1) Our survey is specially effective in the detection of ELGs smaller and less luminous than normal spirals (the median values of the UCM sample are $M_r = -20.6$ and $D_{24} = 19.4$ kpc, for $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$). (2) The sample presents a very heterogeneous morphological population of galaxies. There is a clear predominance of late-type galaxies, with 66% being Sb or later. (3) We have found a set of galaxies specially interesting. In particular, we propose 9 candidates to BCD (blue compact dwarf galaxy), five of them previously unknown. (4) About 10% of the ELGs of the sample show typical parameters of S0 types, and in the subsample of 16 UCM interacting galaxies, 25% are S0 type galaxies. (5) In the set of UCM galaxies inside the Coma cluster, 71% exhibit Hubble types Sb or later.

Key words: surveys — galaxies: fundamental parameters — galaxies: starburst — galaxies: photometry

1. Introduction

The main aims of the UCM $H\alpha$ objective-prism survey are the identification and study of new young, low metallicity galaxies, together with the quantification of the properties of the star formation in the local universe. Details of the UCM survey have already been presented in the list I (Zamorano et al. 1994) and list II (Zamorano et al. 1996) papers. In order to achieve these goals, we have characterized photometrically (Vitores et al. 1996, Paper I of this series) and spectroscopically (Gallego et al. 1996) the UCM sample. From the set of parameters obtained (mainly through surface photometry techniques) in Paper I, in this work we analyse statistically the photometric properties of the UCM sample, and compare the results with those of other samples of galaxies.

In Sect. 2 we describe the statistical properties of the global photometric parameters, i.e. apparent and absolute magnitudes. In Sect. 3 we analyse the bulge+disk

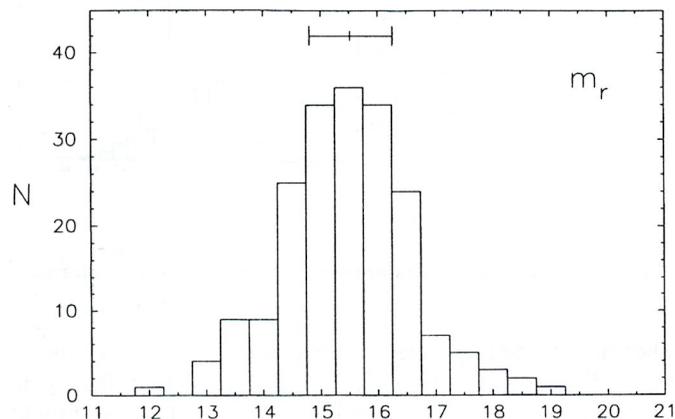


Fig. 1. Frequency histogram of the apparent magnitudes for the UCM sample. In this and the following histograms, the median and quartile points of each distribution are marked

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($B + D$) decomposition parameters, whereas in Sect. 4 results about geometric parameters, mean photometric parameters and concentration indices are presented. Sect. 5 presents correlations and trends among several

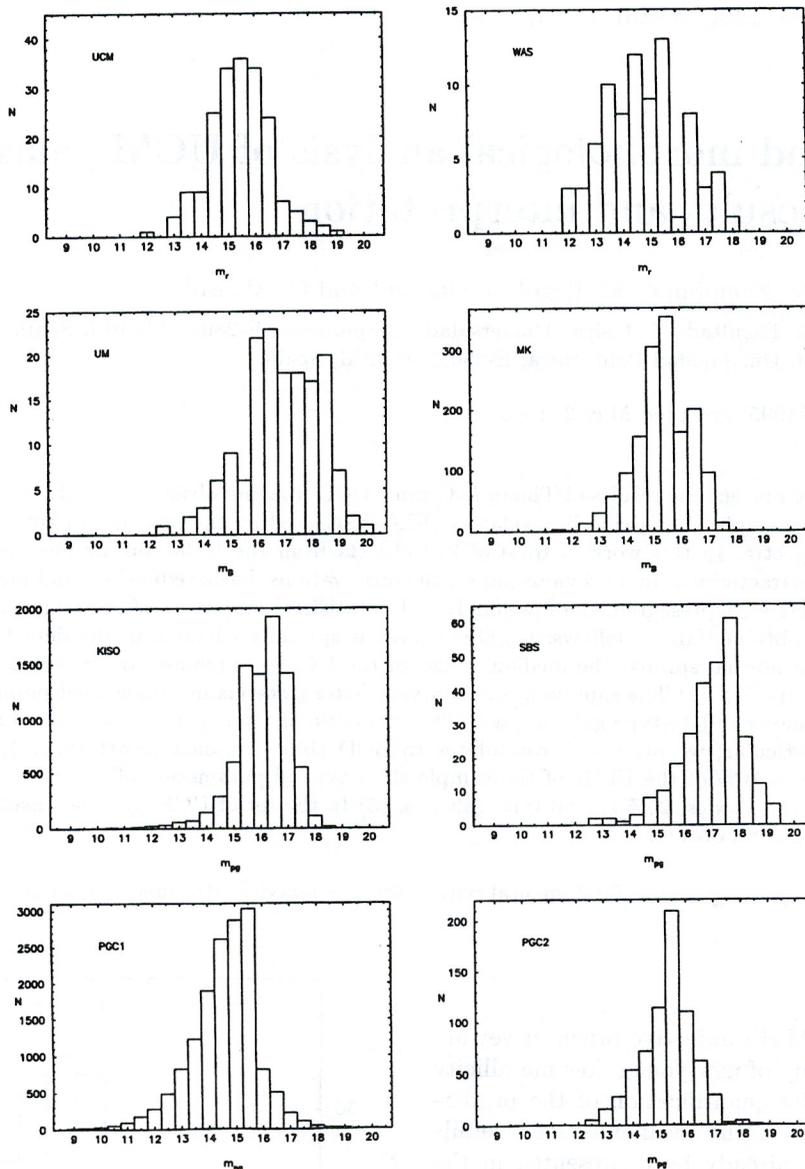


Fig. 2. Histograms of the apparent magnitudes for different samples of galaxies

photometric parameters. In Sect. 6 we perform statistics about the morphology of the UCM objects and, finally, in Sect. 7 we discuss the possible influence of environment on the presence and nature of ELGs.

2. Global photometric parameters

2.1. Distribution of apparent magnitudes

Figure 1 shows the frequency histogram of apparent magnitudes for the UCM objects in the r filter of Thuan & Gunn (1976), m_r , binned in 0.5 mag intervals. The distribution looks symmetrical around a central peak in the 15.5–16.0 interval and the mean value is 15.4. There is

a paucity of objects brighter than $m_r = 13$ since their spectra appear overexposed in the objective-prism plates, with the exception of galaxies with external HII regions of high contrast. There is also a low frequency of objects fainter than $m_r = 18$, due to their low apparent luminosity (unless they present a high H α emission flux). It is important to note that our sample is not apparent magnitude-limited. However, the instrumental objective-prism setup limits the detection to $z \leq 0.045$ (Gallego et al. 1996).

In order to show the influence of the different selection techniques, we compare in Fig. 2 histograms for different samples of galaxies. Wasilewski (WAS) sample (Wasilewski 1983), analysed in the r band by Bothun et al. (1989), presents peaks at $m_r \sim 14.3$ and 15.3 (after

transforming their m_{24} to total m_r using $m_r = m_{24} - 0.21$ according to Kent (1984), and a mean and median values of $m_r = 14.5$, that is, in average the UCM exploration is detecting objects ~ 1 mag fainter than those of the WAS survey.

The distribution of the University of Michigan (UM) survey in the blue (Salzer et al. 1989) is dominated by galaxies in the range $m_B = 16.0$ – 18.5 , with a mean value of $m_B = 16.8 \pm 1.4$ and a median of $m_B = 16.9$, which translates to $m_r = 15.9$ assuming $(m_B - m_r) \simeq 1$ (Kent 1984 and Bothun et al. 1989). Thus the UM survey in the blue, with a higher redshift cutoff, is biased toward fainter objects, detecting galaxies about 0.4 mag (in the median value) fainter than those of the UCM exploration in the red spectral range.

The distribution of the Markarian (MK) exploration (Mazzarella & Balzano 1986), based on a color selection technique (UV excess), shows a more symmetrical shape than that of the UM survey, peaking at $m_B = 15.5$ ($m_r \simeq 14.5$) and showing a low detection of objects fainter than $m_B = 17.0$ ($m_r \simeq 16.0$). The UCM exploration, like the UM survey, gets ~ 2 mag deeper than the MK sample. The comparison with the Kiso survey (Takase & Miyauchi-Isobe 1993 and references therein) is difficult due to the uncertainties for $m_{pg} \geq 16$. Mean magnitude for the Kiso survey is $m_B = 16.1$, showing a tail toward brighter objects. The fraction of objects fainter than $m_B = 17.75$ ($m_r \simeq 16.75$) is only 1.5%, whereas this fraction is 11.3% in the UCM sample.

The Second Byurakan Spectral Sky Survey (SBS) is an objective-prism exploration which makes use simultaneously of three colors. The available apparent magnitude histogram available for this survey is restricted to a subsample of blue compact galaxies (Izotov et al. 1993). This distribution peaks at $m_B \sim 17.7$ and gets, like the UM survey, deeper than the MK exploration. However, the bias of the SBS subsample toward faint galaxies prevent us from further comparisons.

Finally, we have constructed two histograms for normal galaxies using the PGC (Principal Galaxies Catalogue) of Paturel et al. (1989). The histogram named PGC1 corresponds to a subsample of 15330 galaxies with known apparent magnitudes and radial velocities. In comparison with our survey, the PGC1 sample shows a more prominent tail toward brighter objects with a lower fraction of galaxies fainter than $m_r \sim 15.8$ (3% in the PGC1 sample and 11% for the UCM sample). The PGC2 histogram corresponds to a subsample of 705 galaxies selected from the PGC and laying on the UCM exploration fields. When comparing with the UCM sample this distribution is biased (by about 2 mag) toward brighter galaxies.

In summary, the UCM survey is biased against the detection of usually known, bright objects, and it is able to detect fainter objects (by 1–2 mag) than those of color

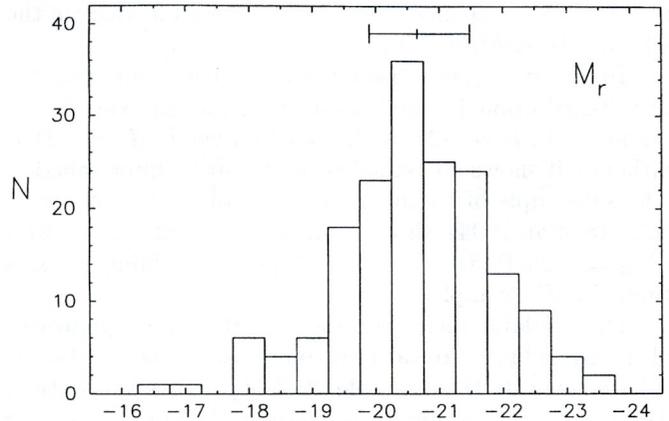


Fig. 3. Histogram of the absolute magnitudes for the UCM sample

surveys. With respect to the blue objective-prism surveys, the UCM exploration selects objects slightly brighter (by 0.4 mag.) than those of the UM survey.

2.2. Distribution of the absolute magnitudes

Absolute magnitudes, M_r , for the objects of the sample have been computed using the radial velocities obtained in the spectroscopic study of Gallego et al. (1996) which uses $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$. In Fig. 3 we show the histogram of the M_r distribution for the UCM sample binned in 0.5 mag intervals. It looks rather symmetrical, peaking at $M_r = -20.5$. Absolute magnitudes span from $M_r = -23.5$ (for the luminous spiral UCM 2317+2356 = NGC 7620 and the cD UCM 2321+1631 = NGC 7647) to $M_r = -16.7$ (for UCM 1612+1308 and UCM 2325+1628, candidates to BCD galaxies). The median of the distribution is -20.65 and the mean magnitude is $M_r = -20.6 \pm 1.2$. Considering that a normal spiral has typically $M_r = -22.0$ (Bothun et al. 1989) or $M_r = -22.5$ (using $M_B = -21.5$ from Simien & de Vaucouleurs (1986) and $B - r \simeq 1$), the UCM survey is detecting, in the mean, galaxies about 2 mag fainter than normal spirals, in agreement with the findings of other ELGs explorations like WAS (Bothun et al. 1989) and UM (Salzer et al. 1989). This confirms the advantage of objective-prism techniques for detecting intrinsically faint objects.

In Fig. 4 we compare the absolute magnitudes histograms for different samples of galaxies (transforming their values to $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ scale). In the WAS sample, analyzed by Bothun et al. (1989), the distribution is more irregular than that of the UCM sample, due to the relatively low number of galaxies (74). The median value of the WAS sample is -20.3 . Concerning the UM sample, the median absolute magnitude is $M_r = -20.0$ ($M_B = -19.0$), i.e. ~ 0.7 mag fainter than the UCM value. The UM survey spans a wider range in luminosities than that

of the UCM, reaching down to $M_r = -13.6$ whereas the UCM ends at $M_r = -16.7$.

To make a comparison with the color surveys, the MK distribution is dominated by bright galaxies, peaking around $M_r \sim -22$, with a median value $M_r = -21.7$, although it shows an extended tail toward fainter galaxies. The subsample of compact galaxies from the Kiso survey (Comte et al. 1994), shows a median value of $M_r = -21.4$ ($M_B = -20.4$), with a lower frequency of faint galaxies than the UM sample.

The available histogram for the SBS survey (Izotov et al. 1993), is biased toward low-luminosity galaxies since it is based on a BCD subsample (25% of galaxies fainter than $M_B = -17.4$). And finally, concerning the normal galaxies catalogues (PGC1 and PGC2, see above), they both peak at $M_r = -22$ ($M_B = -21$), i.e. ~ 1.5 mag brighter than the UCM ELGs sample.

To summarize, surveys based on the presence of emission-lines shows median values of M_r about 1 mag fainter than the samples of galaxies selected by color or the general catalogues. In a relative context, although the UCM exploration presents a contribution of only 8% at the faint range $M_r > -19.0$ (compare with the 31% of the UM blue survey), this contribution is more important than the 2% obtained in the PGC2 subsample.

3. Photometric parameters of the bulge + disk decomposition

The bulge+disk ($B + D$) decomposition of the surface brightness profile for the UCM sample has yield a set of objects (6% of the sample) which do not allow this technique of analysis (see Paper I). However, additional information obtained from the spectroscopic study (Gallego et al. 1996) has shown that all of them are field stars. Thus, in this section we analyse the results from this decomposition technique only for the remaining objects.

3.1. Distribution of the B/D ratios

In Fig. 5, we show the histogram of bulge-to-disk luminosity ratios (B/D) for the UCM survey. As it can be seen, the UCM sample is dominated by low values of B/D : 70% of the galaxies with $B/D \leq 1$ (and 60% with $B/D \leq 0.75$).

From the correlation found by Kent (1985) in the r bandpass between B/D and T (Hubble type), it follows that the UCM sample is dominated by late-type spirals ($T \geq \text{Sb}$), although the B/D ratio is only one of the five morphological parameters used in this work (see Paper I). Moreover, one should be aware of the low reliability of the decomposition for galaxies with $B/D \geq 1.70$ (i.e. $B/T \geq 0.63$). Nevertheless, this is not an important effect in our sample since only 18% of the UCM galaxies exceeds this B/D threshold. When excluding these cases, the mean B/D for the sample is 0.5, a typical value for Sb type galaxies (Kent 1985).

3.2. Distribution of bulge parameters

Figures 6 and 7 represent the histograms for the parameters of the bulge component obtained from the $B + D$ decomposition. In Fig. 6 it is shown that the distribution of the effective bulge surface brightness, μ_e , is dominated by faint bulge components, as it should be expected in a late-type dominated sample of galaxies. The mean μ_e is 22.5 ± 1.6 mag/arcsec² and when excluding the cases with $B/D \geq 1.70$, this value remains almost unalterable (22.7 ± 1.6 mag/arcsec²), indicating that neither the mean value nor the dispersion are an artifact caused by our $B + D$ decomposition technique. Since $\mu_e \simeq 22.0$ for Sb and $\simeq 23.0$ for Sc galaxies (Kent 1985; Simien 1989), it is shown again that the UCM ELGs sample is dominated by Sb-Sc types.

In Fig. 7 we show the histogram of the effective radius, r_e (containing one-half the total light of the bulge component). The sample is again dominated by short scale bulges, with a median value of 2.1 kpc. Only four galaxies exhibit $r_e > 9$ kpc (UCM 1314+2827, with an anomalous $B + D$ decomposition; UCM 2300+2014, with a PSF correction probably overestimated due to a high seeing value; UCM 2320+2428, with a difficult decomposition due to its high inclination, and UCM 2321+1631, an early-type galaxy). Even taking into account the high dispersion in r_e for a given Hubble type (Kent 1985; Simien 1989), the UCM sample is dominated by late-type galaxies (consider that typically $r_e \sim 2.5$ kpc for Sb-Sc). Finally, since the two bulge parameters result more affected by seeing, we must be especially careful when comparing the results between different authors.

3.3. Distribution of disk parameters

Figures 8 and 9 present the histograms for the parameters of the disk component obtained from the $B + D$ decomposition. Figure 8 shows the histogram of the disk (face-on) extrapolated central surface brightness, μ_0 . In contrast to the μ_e histogram, the μ_0 distribution is dominated by bright components, with median and mean values of 20.3 mag/arcsec² (20.2 if we exclude the galaxies with $B/D \geq 1.70$). Moreover, 75% of the sample is concentrated within a narrow range of only 2 mag, in agreement with Kent (1985). There is also a wide range in μ_0 which results narrowed when the cases with $B/D \geq 1.70$ are not considered. Thus, all of the UCM galaxies with $\mu_0 > 23.25$ mag/arcsec² have $B/D > 2.2$ and exhibit morphologic parameters more typical of S0 systems. In these galaxies, as Kent (1985) pointed out, the $B + D$ decomposition is quite unreliable since the reliability of the decomposition depends primarily on the accuracy of $r^{1/4}$ law in describing the bulge, which could produce a spurious disk component.

An important topic about the μ_0 values is the existence of a universal $\mu_0 = 21.65 \pm 0.30$ B-mag/arcsec² as

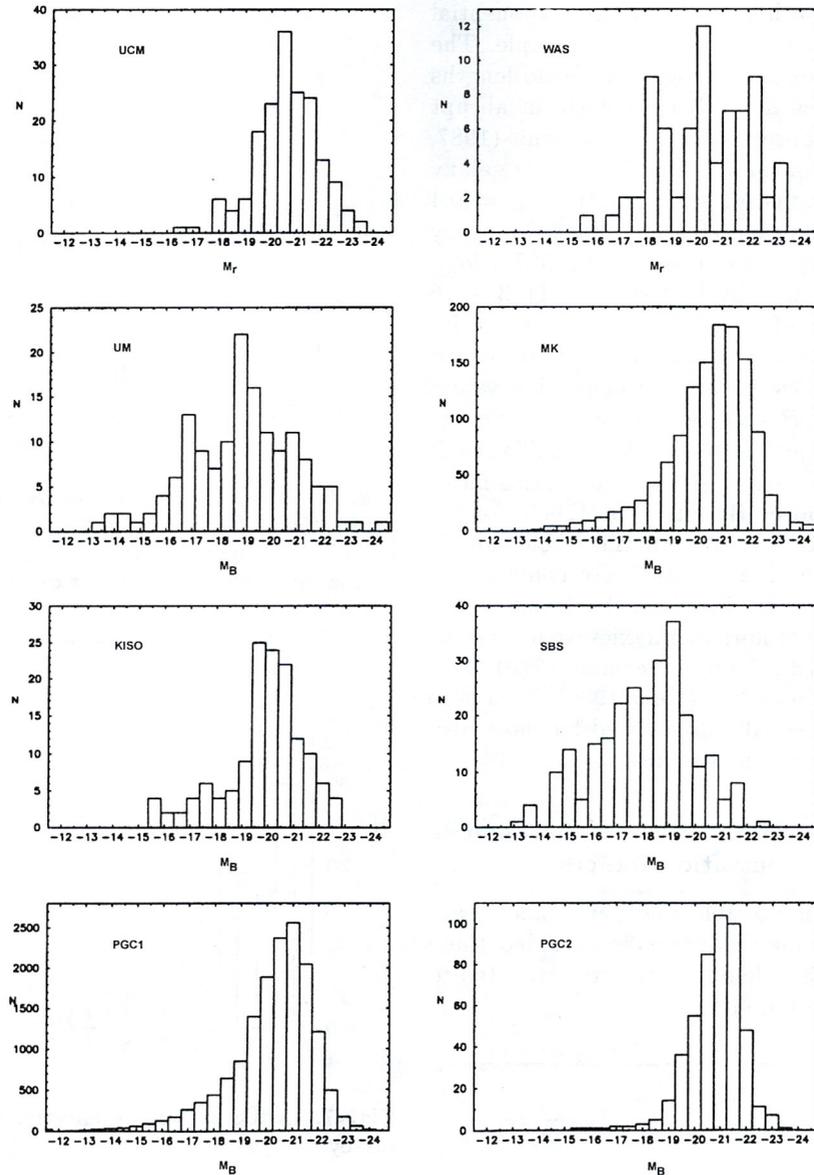


Fig. 4. Histograms of the absolute magnitudes for different samples of galaxies

proposed by Freeman (1970) and confirmed by Boroson (1981) and Simien & de Vaucouleurs (1986). It must be noted, however, that Romanishin et al. (1983) find a mean $\bar{\mu}_0 = 22.74$ B -mag/arcsec² for low surface brightness spirals whilst van der Kruit (1987) finds a brighter value of $\bar{\mu}_0 = 22.5$ B -mag/arcsec² for spirals with disk length of only ~ 1.5 kpc. For the UCM sample, the mean central surface brightness is $\bar{\mu}_0 = 21.3$ B -mag/arcsec² (assuming $B - r \simeq 1$), i.e. 0.35 mag/arcsec² brighter than the Freeman value and 0.2 mag/arcsec² brighter than the mean for the Kent (1985) sample. This result evidences that the disks of ELGs are slightly brighter than those of normal spirals due to the influence of the H α emission (included in the r bandpass). This effect, together with

the trends found for ELGs to be usually smaller and less luminous than normal spirals, has been also observed by Bothun et al. (1989) for the Wasilewski sample. Although their mean value ($\mu_0 \simeq 19.4$ r -mag/arcsec²), is about 1 mag brighter than ours, it must be noted that the bulge component is not considered in the analysis of the surface brightness profiles made by these authors, increasing then artificially the disk contribution. In our sample, we find a high dispersion (1.8 r -mag/arcsec²) around the mean μ_0 value, although this is reduced to 0.9 mag/arcsec² if galaxies with a high B/D ratio are excluded (see above). Thus, the final mean value for μ_0 is 20.2 ± 0.9 mag/arcsec² for the set of UCM galaxies having a more reliable $B + D$ decomposition.

In Fig. 9 we present a histogram for the exponential scale length of the disk, d_L , for the UCM sample. The distribution is clearly dominated by short scale lengths (65% of the sample show $d_L < 3$ kpc) with an abrupt fall for higher values, as observed by van der Kruit (1987, 1989). Excluding the artificial d_L value for the cD galaxy UCM 2321+1631, our range of scale lengths ($d_L = 0.4 - 14.7$ kpc) is intermediate between the ranges found by Freeman (1970), 1 – 12 kpc, or van der Kruit (1987, 1989), 1 – 11 kpc, and those found by Boroson (1981), 3 – 16 kpc, or Kent (1985), 1.5 – 0 kpc. The mean d_L value for the whole UCM sample is 3.2 kpc, although with a high dispersion of 2.8 kpc. These mean and dispersion values are not produced by our $B + D$ decomposition technique since they remain unchanged even if the less reliable $B + D$ decomposition galaxies are excluded. Moreover, this mean d_L is similar to that found by Bothun et al. (1989) for the Wasilewski sample of ELGs (4 kpc for the whole sample and 3.2 kpc for galaxies with $z < 0.033$). For comparison, our Galaxy presents $d_L = 7.5$ kpc (van der Kruit 1989), and from several samples of normal galaxies we have computed a mean $d_L = 4.1 \pm 2.6$ kpc (Freeman 1970), 7 ± 3 kpc (Boroson 1981) and 6 ± 4 kpc (Kent 1985). Therefore the UCM sample is statistically dominated by short disk scales, more typical of late-type galaxies.

4. Geometric parameters, mean photometric parameters and concentration indices

Since even for the stars is possible to obtain some meaningless photometric parameters, we have excluded these objects from any analysis. Therefore, we are restricted to a set of 194 galaxies (see Paper I).

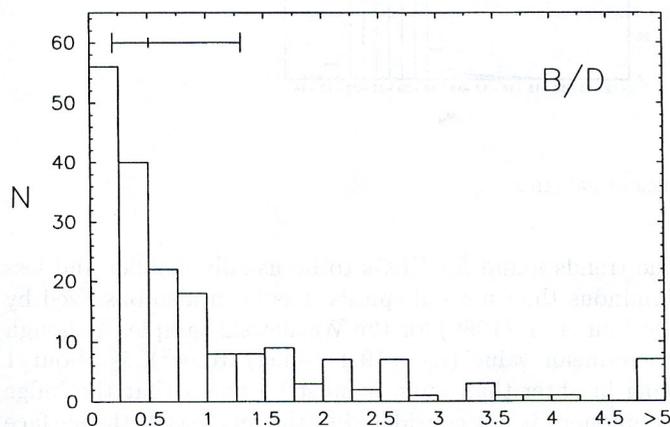


Fig. 5. Histogram of the B/D ratios for the UCM survey

4.1. Geometric parameters

Figure 10 presents a histogram of the diameters, measured at the 24 mag/arcsec² level, for the UCM galaxies. There is

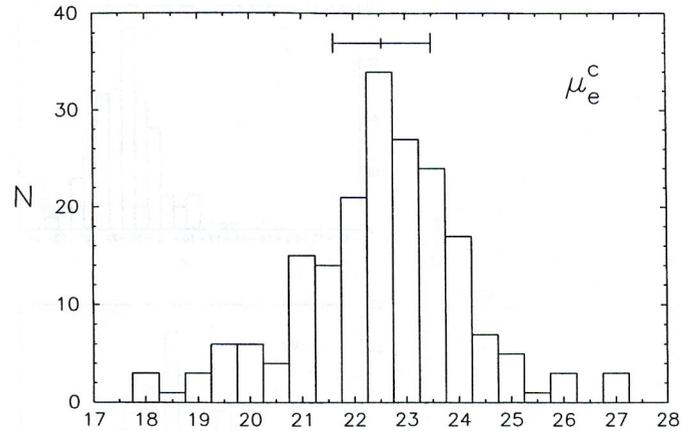


Fig. 6. Histogram of the effective bulge surface brightness, μ_e , for the UCM survey (in this and the following figures, superindex “c” denotes corrected for galactic extinction and for inclination, when necessary, as exposed in Paper I)

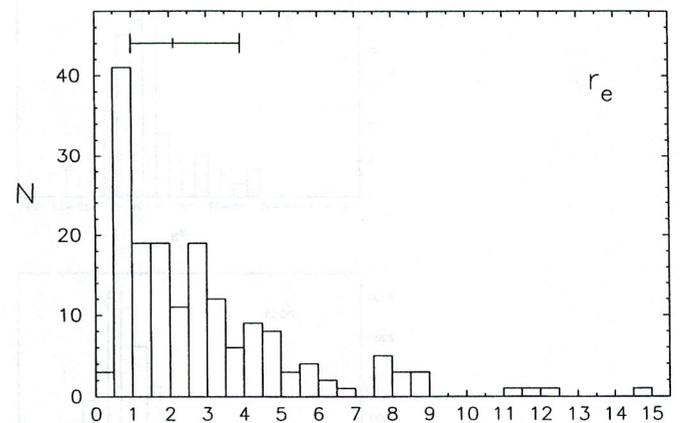


Fig. 7. Histogram of the effective radius, r_e , for the UCM survey

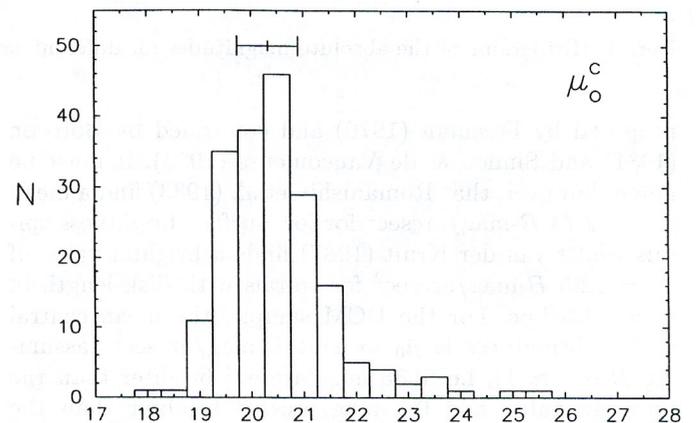


Fig. 8. Histogram of the disk face-on extrapolated central surface brightness, μ_0 for the UCM survey

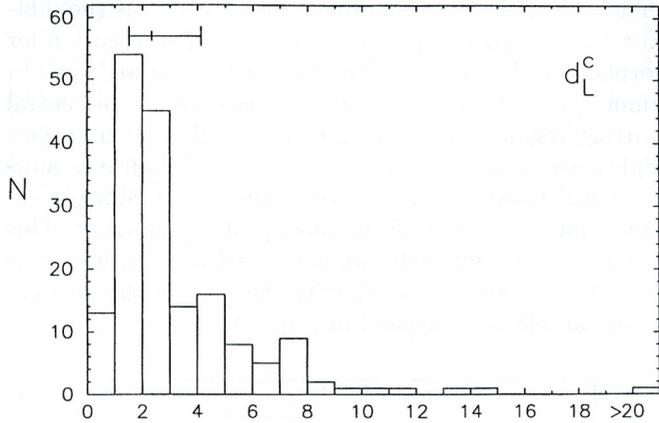


Fig. 9. Histogram of the exponential scale length of the disk, d_L^c , for the UCM survey

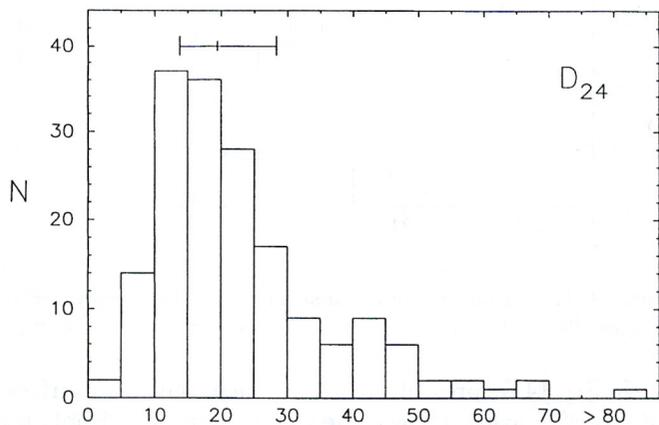


Fig. 10. Histogram of the diameters, measured at the 24 mag/arcsec² level, for the UCM survey

a clear predominance of the small sizes, with 78% of galaxies with $D_{24} < 30$ kpc (nearly the size of our Galaxy). This is an expected result since the objective-prism technique favours the detection of smaller and less luminous systems in which emission features appear more contrasted, as Bothun et al. (1989) pointed out in their study of the Wasilewski survey. Our sample covers a wide range of sizes (3 – 67 kpc, excluding the cD UCM 2321+1631), the median is 19.4 kpc and the mean value is 23 ± 14 kpc.

For comparison, the Wasilewski sample (Bothun et al. 1989) ranges up to 90 kpc, being dominated by the smallest galaxies. Although their fraction of galaxies larger than 30 kpc is 36% (higher than in our sample), their whole sample has a mean value of 26 kpc, quite similar to our result. The available data for the University of Michigan survey (Salzer et al. 1989) have allowed us to compare our D_{24} values in the r band with D_{25} in the B band. The UM sample extends up to 80 kpc and exhibits also a predominance of small galaxies, with 29% having $D_{24} < 7.5$ kpc

(5% in the UCM sample) and only 10% of galaxies larger than 30 kpc (22% in the UCM sample). The mean diameter for the UM is 17 kpc, although with a large dispersion of ± 14 kpc, and with a median value of 14 kpc, i.e. $\sim 30\%$ smaller than in our sample. In summary, the UCM survey is detecting, in average, galaxies slightly smaller than the WAS survey but larger than the UM sample. Finally, it must be noted that size is one of the most important parameters to classify ELGs into different groups (Salzer et al. 1989). These authors find that all the galaxies with diameters above 22.5 kpc are Sy or SBN (starburst nuclei) and, using a sample of 48 Sy from MacKenty (1990), we compute a mean $\bar{D}_{24} = 25 \pm 10$ kpc. In the UCM sample, we confirm the trend of these active galaxies to have larger sizes ($\bar{D}_{24} = 36 \pm 15$ kpc for our subset of Sy galaxies).

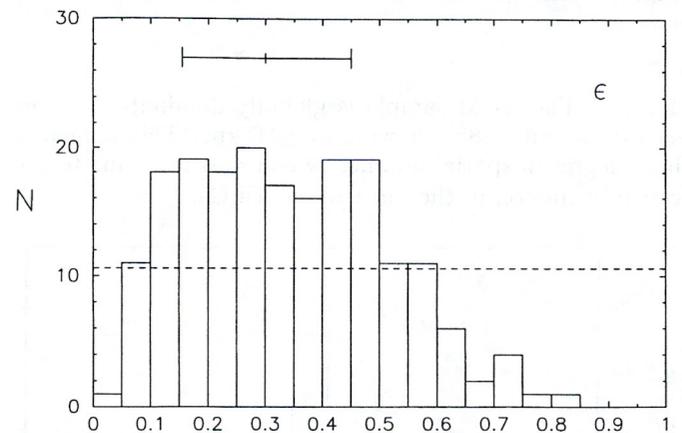


Fig. 11. Histogram of the observed ellipticity ϵ for the UCM survey. The dashed line indicates the frequency for randomly orientated disks

Another geometric parameter that we have calculated in our sample is the ellipticity ϵ (averaged over the five most external isophotes better defined). The distribution, shown in Fig. 11, reveals a nearly uniform distribution for $\epsilon < 0.5$ and a steep decline of detections above this value. The distribution does not show the expected uniformity for a sample of disks randomly oriented, but is dominated by low-inclinations objects. Thus, only 17% have $\epsilon > 0.5$, showing the difficulty in detecting nearly edge-on systems due to their higher internal extinction.

4.2. Mean photometric parameters

In Figs. 12 to 14 we present the histograms of the mean photometric parameters, calculated in Paper I. These parameters are important due to their independence on the decomposition procedure used in the analysis of the radial distribution of the light in the galaxies. Figure 12 shows the histogram of the effective radius, a_e (containing one-half the total luminosity of a galaxy). The median is 2.8 kpc and the mean is 3.8 kpc, but with a dispersion of

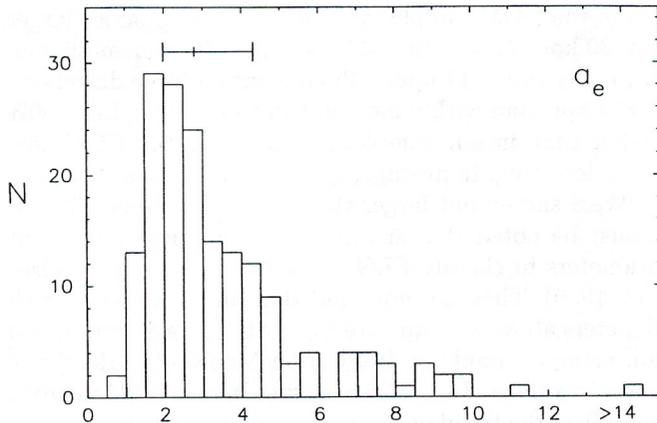


Fig. 12. Histogram of the effective radius, a_e , for the UCM survey

4.6 kpc. The UCM sample is globally dominated by low a_e values, with 88% showing $a_e \leq 6$ kpc. This reflects a high degree of spatial luminosity concentration due to the central emission in the most of our ELGs.

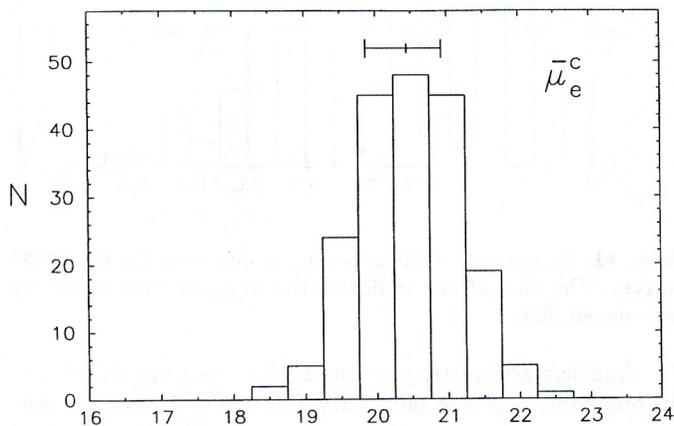


Fig. 13. Histogram of the mean surface brightness inside the effective radius, $\bar{\mu}_e$, for the UCM survey

In Fig. 13, the distribution of mean surface brightness inside the effective radius, $\bar{\mu}_e$, is shown. Although the histogram covers a range of about 4.5 mag, quite similar to that found by Kent (1985), most of the galaxies spans over a range of only 1.5 mag (19.75 – 21.25 mag/arcsec²) which includes 67% of the sample, and the mean value is 20.46 ± 0.70 mag/arcsec². In comparison, we have obtained from the Kent (1985) sample a mean value of 20.73 ± 0.70 mag/arcsec² for normal galaxies, whereas for the low surface brightness galaxies (LSB) from Romanishin et al. (1983) the mean is ~ 22.6 mag/arcsec² (using $B - r \simeq 1$). Therefore, on the one hand, the mean value of $\bar{\mu}_e$ for the UCM sample is ~ 0.3 mag brighter than that of normal galaxies. And, on the other hand, the relatively low dis-

persion of the values (0.70 mag/arcsec²) reveals the difficulty in using $\bar{\mu}_e$ as the only parameter of segregation for morphological types, as Kent (1985) has pointed out. In summary, in our sample of ELGs the presence of central emitting regions causes an increase in the mean surface brightnesses inside the effective radius and its use as morphological parameter is not very convenient, since it excludes one-half the total luminosity of the galaxies. This is one of the reasons why we have used a set of five morphological parameters to classify the type of the galaxies in our sample, as discussed in Paper I.

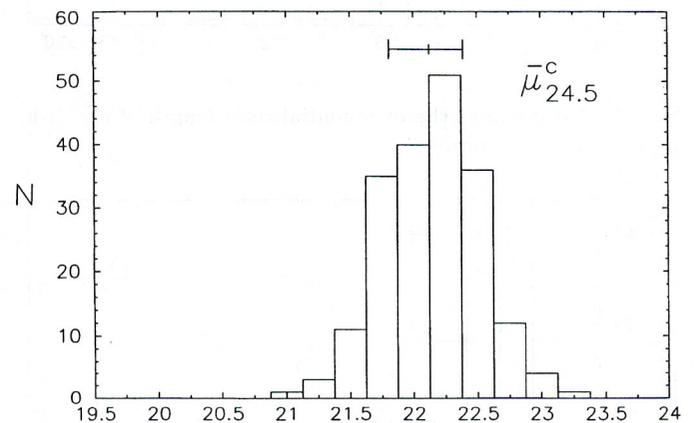


Fig. 14. Histogram of the mean surface brightnesses averaged over the 24.5 mag/arcsec² isophote, $\bar{\mu}_{24.5}$, for the UCM survey

In Fig. 14 we present the histogram of the mean surface brightnesses averaged over the 24.5 mag/arcsec² isophote. The range for ELGs spans only about 2.5 mag and 78% of the sample is inside a narrow range of only 1 mag (21.6 – 22.6 mag/arcsec²). Considering the diversified morphology of the UCM galaxies, that result reflects the insufficiency of using a mean surface brightness as the only parameter for morphological segregation. The mean value is 22.14 ± 0.36 mag/arcsec², typical of normal galaxies (Doi et al. 1993; Okamura et al. 1994). Thus, whereas $\bar{\mu}_e$ is quite sensitive to the presence of nuclear bursts, $\bar{\mu}_{24.5}$, averaged over almost the whole galaxy, is unaffected by the emission features. Moreover, the low dispersion on $\bar{\mu}_{24.5}$ indicates that this mean brightness can not be used as the only parameter for morphological segregation.

4.3. Concentration indices

Histograms for the three concentration indices used in Paper I are shown in Figs. 15-17. Figure 15 represents the distribution for $c_{in}(\alpha)$, with $\alpha = 0.3$, a parameter introduced by Doi et al. (1993). The range is $0.24 \leq c_{in}(\alpha) \leq 0.70$, similar to that found by Okamura et al. (1994) for normal galaxies in the R band. The distribution is quite symmetrical around a median of 0.48, and with 66% of the objects in the range 0.40–0.60. The mean

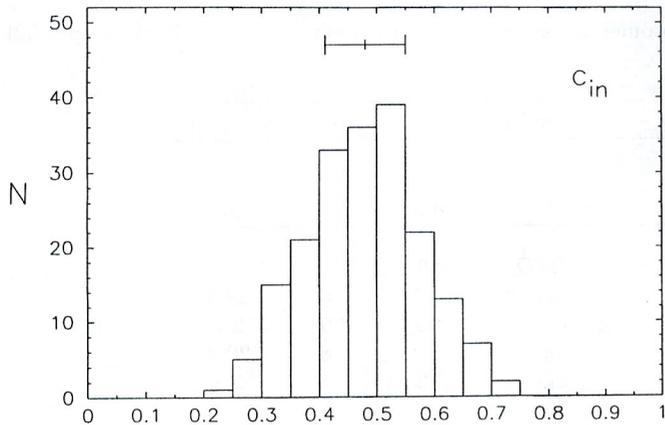


Fig. 15. Histogram for the $c_{in}(\alpha)$ values for the UCM survey

value for the UCM ELGs is 0.47 ± 0.10 , a typical result for Hubble types later than Sa (Doi et al. 1993). Then, due to the insufficiency of this index as the only morphological parameter, we have used the biparametrical classification ($c_{in}(\alpha)$, $\bar{\mu}_{24.5}$) proposed by these authors as one of the five morphological criteria (see Paper I).

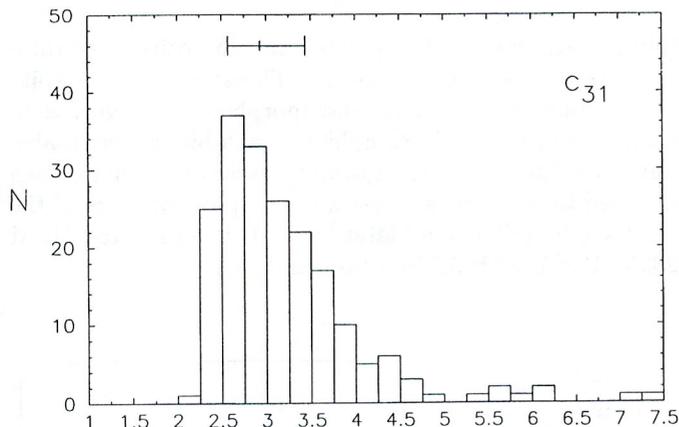


Fig. 16. Histogram for the values of the c_{31} concentration index for the UCM survey

In Fig. 16 we present the histogram for the c_{31} concentration index introduced by de Vaucouleurs (1977). The wide range reveals the diversity of morphological types in our ELG sample (2.8 is the theoretical value for a pure exponential profile and 7.0 for a pure $r^{1/4}$ law). The histogram is mainly populated by low values of c_{31} , as expected for a sample dominated by late-type galaxies. The median is 3.0 and the mean is 3.2 ± 0.8 , a value corresponding to Hubble types later than Sa (Gavazzi et al. 1990), and 83% of the values fall between 2.25 and 3.75. A steep decline is observed for $c_{31} < 2.25$, a value only attained by some Irr in the sample of Gavazzi et al. (1990). In fact, in the UCM sample, only UCM 1303+2908 (an irregu-

lar ELG in Coma) is below this value. The deficiency of galaxies with $c_{31} > 4.5$ represents the paucity of Hubble types S0 or earlier in our sample. Moreover, 42% of the sample shows $c_{31} < 2.8$ (the theoretical value for a pure disk model), a similar value to the 49% found from the sample of normal galaxies of Gavazzi et al. (1990). Concerning the Sy subsample, although some of them exhibit c_{31} as high as 7.21, others have values of 2.65. Thus, as expected, this parameter does not represent any advantage in segregating active galaxies. Finally, for the UCM subsample of BCDs, we obtain for c_{31} a mean of 2.6 ± 0.3 , that is a very low value when comparing with the mean for the whole sample, suggesting a rather late Hubble type for this kind of galaxies.

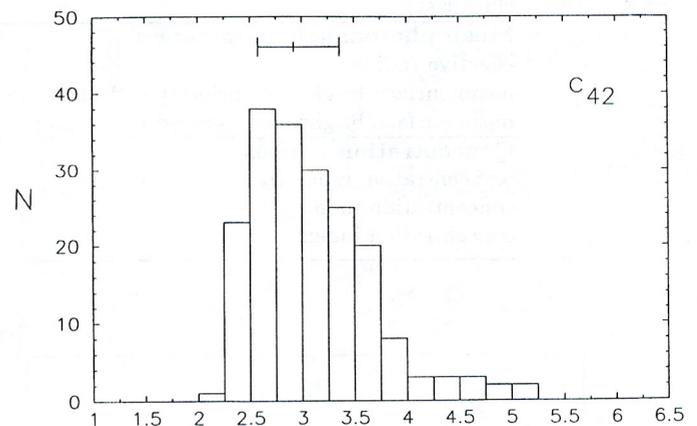


Fig. 17. Histogram for the values of the c_{42} concentration index for the UCM survey

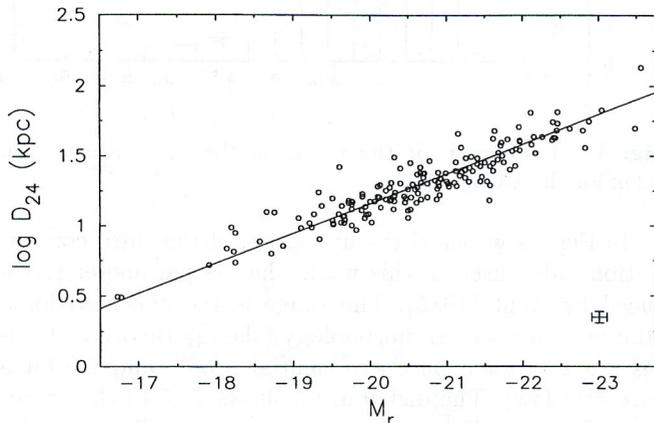
In Fig. 17 we show the histogram of the third concentration index used in this work, the c_{42} parameter introduced by Kent (1985). The range is the expected for a sample with a varied morphology (the c_{42} theoretical values are 2.74 for a pure exponential profile and 7.0 for a pure $r^{1/4}$ law). The distribution shows 78% of the values in the 2.25 – 3.75 band, a typical range for Sb morphologies or later (Kent 1985). The median value is 3.0 and the mean is 3.1 ± 0.6 , a typical value for Sb–Sc⁺ types. The relatively low fraction of galaxies with $c_{42} = 4.50$ – 5.13 is dominated by galaxies that correspond to S0 types. And, finally, concerning the Sy and BCD subsamples, the results are the same as those for the c_{31} index.

5. Correlations between the absolute, geometrical, mean photometric and concentration parameters

In the previous section we have presented the statistical results of the UCM sample. These results have been summarized in Table 1 for a quick look. Now, in this section we analyse some trends and correlations between the parameters discussed above.

Table 1. Mean, standard deviation and median values of the photometric parameters for the galaxies in the UCM sample (all the scales are in kpc and the surface brightnesses in mag/arcsec²)

	symbol	mean	st.dev.	median
Magnitudes				
apparent magnitude	m_r	15.4	1.1	15.4
absolute magnitude	M_r	-20.6	1.2	-20.7
B + D parameters				
bulge to disk ratio	B/D	0.9	1.1	0.5
effective bulge surface brightness	μ_e	22.5	1.6	22.5
effective radius of the bulge	r_e	3.0	3.0	2.1
disk face-on central surface brightness	μ_0	20.3	1.8	20.3
exponential scale length of the disk	d_L	3.2	2.8	2.3
Geometric parameters				
diameters at 24 mag/arcsec ²	D_{24}	23.0	13.6	19.4
ellipticity	ϵ	0.3	0.2	0.3
Mean photometric parameters				
effective radius	a_e	3.8	4.6	2.8
mean surface brightness inside the effective radius	$\bar{\mu}_e$	20.5	0.7	20.5
mean surface brightness averaged over 24.5 mag/arcsec ²	$\bar{\mu}_{24.5}$	22.1	0.4	22.1
Concentration indices				
concentration index ($\alpha = 0.3$)	$c_{in}(\alpha)$	0.5	0.1	0.5
concentration index	c_{31}	3.2	0.8	3.0
concentration index	c_{42}	3.1	0.6	3.0

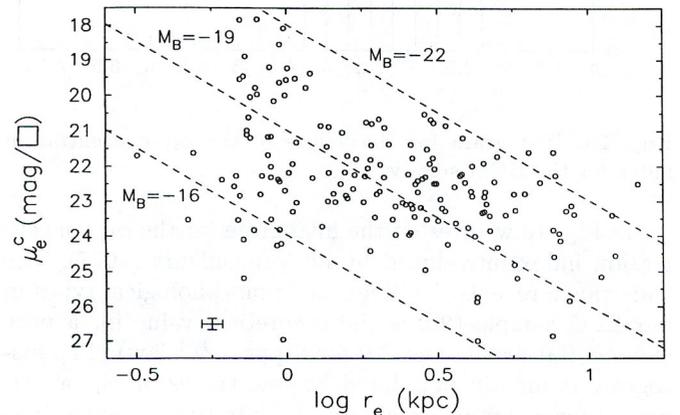
**Fig. 18.** Correlation between the diameter at the 24 mag/arcsec² level and the absolute r magnitude for the UCM survey. A linear least-squares fit to the data is shown by a straight line. In this and the following plots, typical error bars are marked

5.1. Size-luminosity relation

Figure 18 shows the tight correlation between the diameter at the 24 mag/arcsec² level and the r absolute magnitude for the UCM sample. A least-square fit to our data leads to:

$$\log D_{24} = (-3.10 \pm 0.13) + (-0.21 \pm 0.01)M_r \quad (1)$$

(Note that, analytically, a constant luminosity/area ratio predicts $\log D = C - 0.20M$). Therefore, and in spite of the different processes and morphologies involved in the ELGs of the UCM sample, this exhibits a remarkable uniformity in mean surface brightnesses. This trend is even followed by the two weakest and compact galaxies of the sample (the BCD candidates UCM 1612+1308 and UCM 2325+1628), with $M_r \sim -16.7$ and $D_{24} \sim 3$ kpc.

**Fig. 19.** μ_e vs. $\log r_e$ distribution for the bulge component of the UCM survey. Lines of constant absolute magnitude for bulges are indicated by dashed lines

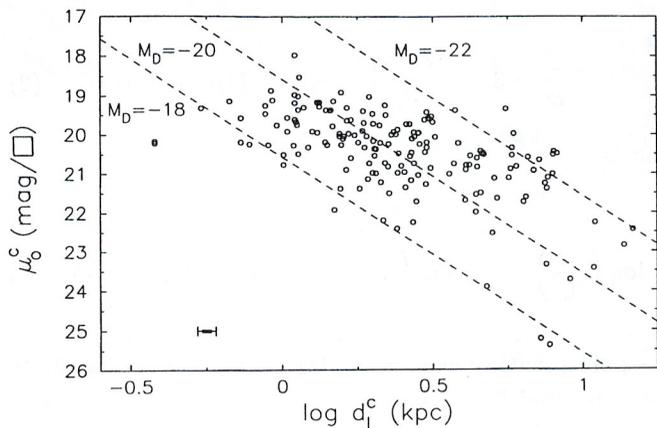


Fig. 20. μ_0 vs. $\log d_L$ distribution for the disk component of the UCM survey. Lines of constant absolute magnitude for disks are indicated by dashed lines

5.2. Relations concerning $B+D$ parameters

Figures 19 and 20 show the μ_e vs. $\log r_e$ distribution for the bulge component and the μ_0 vs. $\log d_L$ for the disk component, respectively. Concerning the bulge distribution, there is some evidence of a trend between radius and surface brightness. Comparing with the lines of constant luminosity ($M_{\text{Bulge}} = \text{constant}$ corresponds to $\mu_e = C + 5 \log r_e$), the large scatter of points indicates a wide range in absolute magnitudes. The same effect is observed by Kent (1985), Simien & de Vaucouleurs (1986) and Kodaira et al. (1986), although the effect of the seeing correction on the bulge parameters prevent us from further comparisons. A similar trend between the disk parameters is observed in Fig. 20 ($M_{\text{Disk}} = \text{constant}$ corresponds to $\mu_0 = C + 5 \log d_L$). In this case, the scatter is lower than that obtained for the bulges and it is similar to that found by Kodaira et al. (1986), although a bit larger than that observed by Simien & de Vaucouleurs (1986). Moreover, this diagram shows that the majority of the disks in our sample exhibit disk luminosities weaker than those of normal late type spirals ($M_{\text{Disk}} \sim -22$). Since μ_0 value is (in the mean) slightly brighter than in normal spirals, that effect is explained as due to shorter disk length-scales in our ELGs.

Also in connection with the $B+D$ decomposition, we have investigated the trends between the four parameters derived from this technique and the bulge-to-total luminosity ratio, B/T , as proposed by Kent (1985). In principle, a drop in the B/D ratio may be due to either a decrease in bulge size relative to the disk, a drop in bulge surface brightness, or a combination of both effects (Kent 1985). To investigate this, on the one hand, we plot in Fig. 21 the relation between the ratio of scale lengths and B/T . Although there is a trend in the sense that the scale lengths ratio increases with B/T (especially if the less reliable cases, with $B/T \geq 0.63$, are excluded), the scatter on

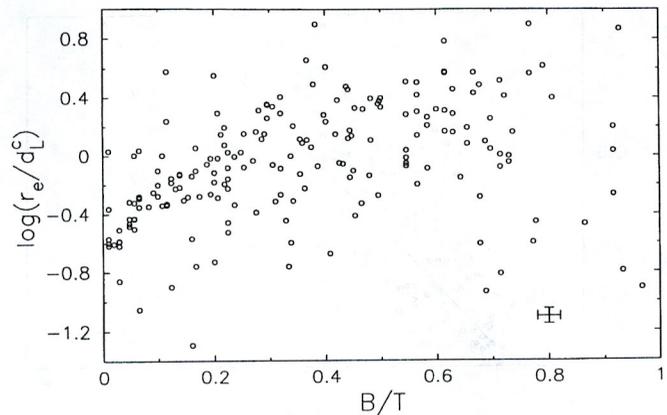


Fig. 21. Relation between the ratio of scale lengths and the B/T ratio for the UCM survey

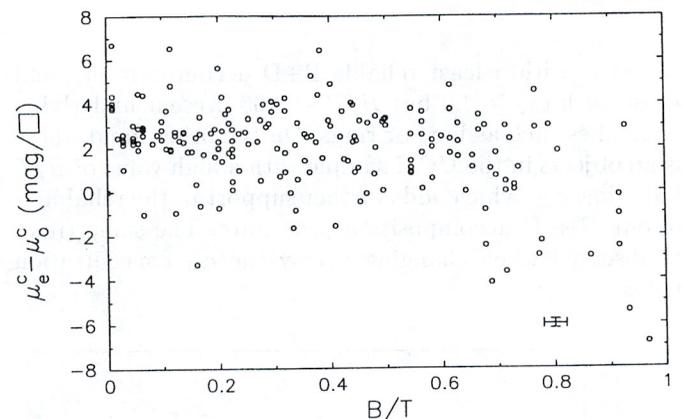


Fig. 22. Relation between the difference of surface brightnesses and the B/T ratio for the UCM survey

the plot is remarkable. On the other hand, although this author proposes that a dimming of the bulge rather than a brightening of the disk would be primarily responsible for a drop in B/T , for our sample there is no relation between the difference of surface brightnesses and B/T ratio (Fig. 22). However, the difficulty in unifying the different $B+D$ decomposition techniques could be responsible for the observed discrepancies among different authors.

Since the concentration indices are introduced due to their advantage in measuring the concentration of light in the galaxies with independence of any decomposition procedure, the relation between B/T ratio and the concentration index c_{42} has been investigated. In Fig. 23 we observe a trend in B/T to increase with c_{42} , as found by Kent (1985). Only objects with $c_{42} \geq 4.0$ (corresponding to early type galaxies) deviate from the mean relation, but since 87% of the UCM sample has $c_{42} < 3.75$, the contribution of those points is negligible. On the one hand, all the cases with $B/T > 0.63$ and $c_{42} \geq 4.0$ correspond to S0

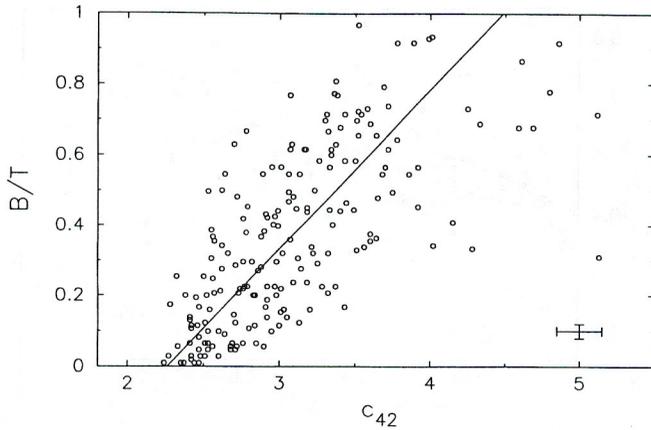


Fig. 23. Relation between the B/T ratio and the concentration index c_{42} for the UCM survey. A linear least-squares fit to the data is shown

galaxies (with a least reliable $B+D$ decomposition), and those with $c_{42} \geq 4.0$ but $B/T < 0.63$ present underlying anomalies such as bars or rings. On the other hand, there is no objects in the UCM sample with a high value of B/T but a low c_{42} , which adds further support to the reliability of our $B+D$ decomposition procedure. The same trend is observed when changing c_{42} by the c_{31} concentration index.

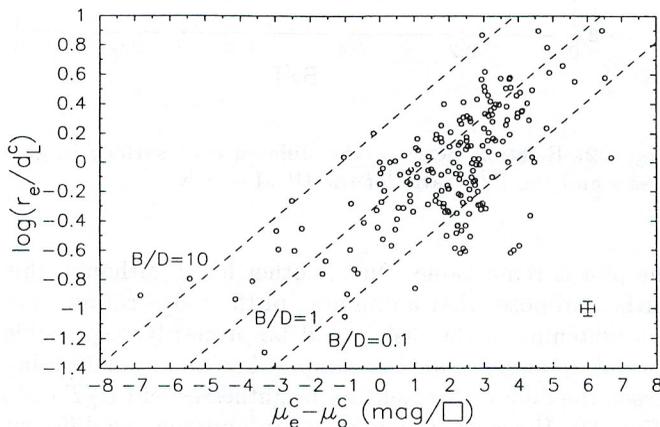


Fig. 24. Dependence between the ratio of scale lengths and the difference of surface brightnesses for the UCM survey. The dashed lines show the positions of constant B/D ratio in sequences of $\log(B/D) = \pm 1$

We have also analysed the dependence between the ratio of scale lengths and the difference between bulge and disk surface brightnesses. The trend, represented in Fig. 24, is similar to that observed by Kodaira et al. (1986) for a heterogeneous sample of 167 galaxies. This plot corresponds in fact to a distribution of the objects according to their B/D ratio, since from

$$\frac{B}{D} = 3.607 \left(\frac{r_e}{d_L} \right)^2 \text{dex}(-0.4(\mu_e - \mu_0)) \quad (2)$$

that is

$$\log \left(\frac{r_e}{d_L} \right) = 0.5 \log \left(\frac{B}{D} \right) - 0.279 + 0.2(\mu_e - \mu_0) \quad (3)$$

the points with a constant B/D ratio would distribute along straight lines with

$$\log \left(\frac{r_e}{d_L} \right) = C + 0.2(\mu_e - \mu_0) \quad (4)$$

Therefore, and as can be seen in the figure, most of the UCM objects are below $B/D = 1$, as corresponds to a late-type dominated sample, and the region with $B/D > 10$ is populated by S0 types.

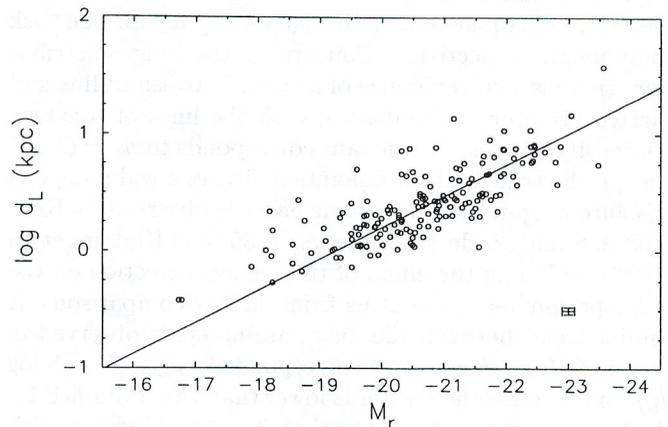


Fig. 25. Correlation between M_r and the exponential scale length of the disk, d_L , for the UCM survey. A linear least-squares fit to the data is shown

Finally, we have investigated the correlations of the parameters obtained from the $B+D$ decomposition with the absolute magnitude of the galaxies. The only parameter that shows a clear correlation with M_r is the disk exponential scale length. This dependence, shown in Fig. 25, represents a clear influence of disk scales in the total luminosity of galaxies, as expected for disk-dominated systems. A linear least-squares fit to the data leads to the following relation:

$$\log d_L = (-5.20 \pm 0.29) + (-0.27 \pm 0.01)M_r \quad (5)$$

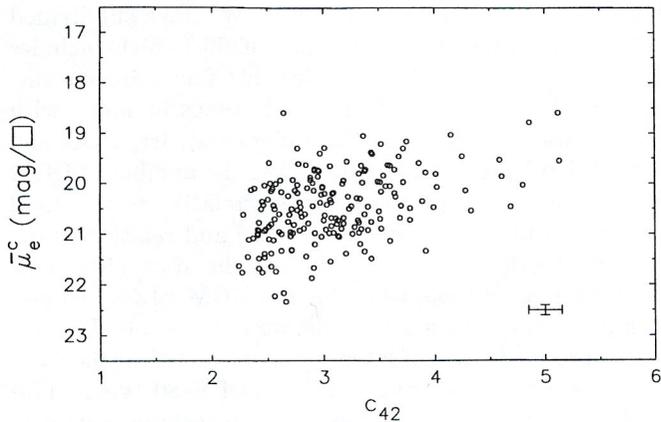


Fig. 26. Mean surface effective brightnesses $\bar{\mu}_e$ as a function of the c_{42} concentration index for the UCM survey

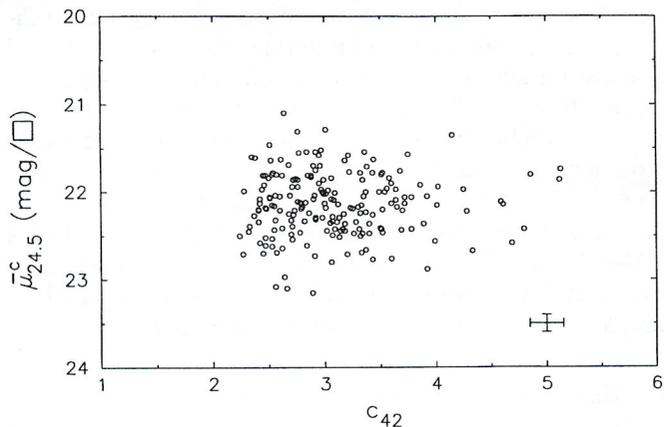


Fig. 27. Mean surface brightness inside the 24.5 mag/arcsec² isophote, $\bar{\mu}_{24.5}$, as a function of the c_{42} concentration index for the UCM survey

5.3. Relations concerning mean photometric parameters

We have studied the relations between mean surface brightnesses and concentration indices. In Fig. 26 a trend is shown in the sense that high c_{42} objects (early-type galaxies) exhibit relatively brighter mean surface effective brightnesses $\bar{\mu}_e$, in agreement with Kent (1985). The same effect is obtained when replacing c_{42} by c_{31} . This trend reflects that both parameters represent a good quantitative estimation of the concentration of the light in the galaxies. However, the mean surface brightness inside the 24.5 mag/arcsec² isophote, $\bar{\mu}_{24.5}$, is rather independent of the c_{42} or c_{31} indices, as shown in Fig. 27, since the concentration of the light is diluted when averaging over the entire galaxy.

Although $\bar{\mu}_{24.5}$ is insensitive to the Hubble type, Doi et al. (1993) showed its utility as a segregation morphological parameter when combining it with the $c_{in}(\alpha)$ in-

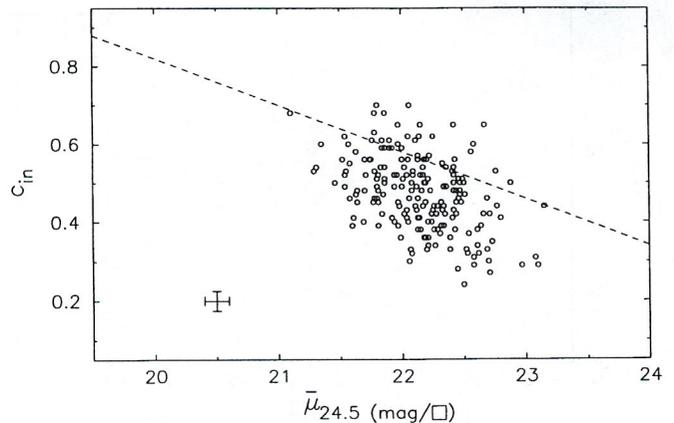


Fig. 28. Positions on the (c_{in} , $\bar{\mu}_{24.5}$) plane of Doi et al. (1993) for UCM objects. The dashed line is the segregation line between early and late-type galaxies proposed by these authors

dex. In fact, the plane defined by these two parameters has been one of the five morphological classifiers used in Paper I. Figure 28 shows the distribution of the UCM sample in this plane. The majority of the UCM objects are placed below the straight line of segregation proposed by Doi et al. (1993), that is, corresponds to Hubble types Sa or later. The UCM sample is confined within a region of about 2 mag in $\bar{\mu}_{24.5}$ and 0.5 units in $c_{in}(\alpha)$. All of the objects above the segregation line have a high B/D ratio, corresponding to early types. In particular, all the Sy 1 in our sample are within the region $\bar{\mu}_{24.5} = 21.8 - 22.4$ mag/arcsec² and $c_{in}(\alpha) = 0.6 - 0.7$, whereas the Sy 2 lay below that straight line and scatter across all the morphologies. This fact confirms the known trend of the Sy 1 nuclei to be hosted in early types galaxies, whereas Sy 2 are usually located in later types. Moreover, all the UCM galaxies with an aspect clearly irregular on the direct image and classified also as Irr according to our five-criteria method (Paper I), are located in the plot towards the lower right corner, as expected for the latest Hubble types, and the same trend is observed for the BCD candidates.

Finally, in Fig. 29 we have analysed the possible dependence between $\bar{\mu}_{24.5}$ and the absolute magnitude M_r . The observed independence of $\bar{\mu}_{24.5}$ on the total luminosity agrees with a roughly constant light/area ratio when averaging over the entire galaxy.

6. Morphological statistics of the objects in the UCM survey

In this section we present a statistical analysis of the morphological aspects (location of emitting knots, Hubble types, tidal features, etc.) for the UCM objects.

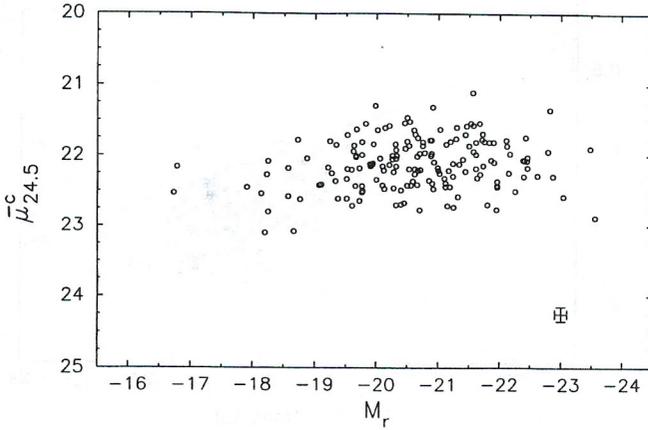


Fig. 29. Mean surface brightness inside the 24.5 mag/arcsec² isophote, $\bar{\mu}_{24.5}$, as a function of the absolute magnitude M_r for the UCM survey

6.1. Location of emitting knots. Frequency of Hubble types

We have found that only 22% of the UCM galaxies (42/194) show star formation regions outside the nucleus. In no way this result can be extrapolated to the relative frequency of ELGs with non-nuclear star forming regions in the Local Universe, since our survey, like others, presents complicated selection effects.

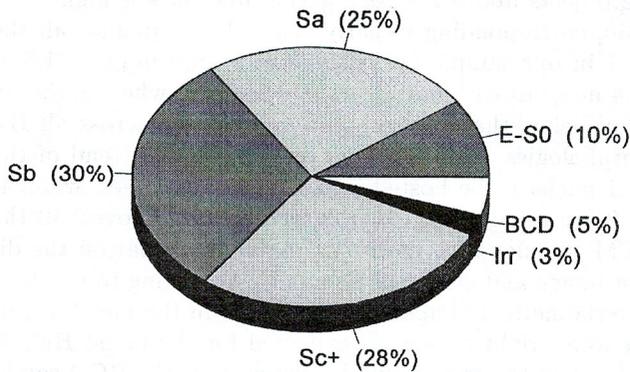


Fig. 30. Diagram of the distribution of Hubble types for the UCM ELGs

Table 2. Hubble types for the UCM objects

Type	E-S0	Sa	Sb	Sc ⁺	Irr	BCD	Stars
<i>N</i>	19	48	59	54	5	9	18
objects	9%	23%	28%	25%	2%	4%	8%
only galaxies	10%	25%	30%	28%	3%	5%	—

Concerning the morphologies, we have carried out a statistical analysis of the UCM sample and also in compar-

ison with other samples of galaxies. We have subdivided the sample in the following types: E-S0 (which includes all the UCM S0 and the elliptical cD found in our survey), Sa, Sb, Sc⁺ (that includes Sc types or later, with the exception of Irr and BCD categories), Irr, BCD and stars. Table 2 presents, for each type, the number of UCM objects and the respective fractions relative to the total number of objects in the survey (212) and relative to the number of galaxies (194), excluding the stars (18). This distribution of Hubble types for the UCM ELGs has also been represented by a sector diagram in Fig. 30. There is a clear predominance of later types, with 66% of the sample being Sb or later and only 10% of E-S0 types. This result is a consequence of the selection technique used in the UCM survey, since the galaxies have been selected by the presence of emission features. Within the latest types, the paucity of Irr galaxies in the UCM survey could come from the fact that, in these systems, the emitting regions are not contrasted enough to be detected by our technique. The predominance of late types in samples selected by emission features has been already observed by many authors (Gisler 1978; and Moss & Whittle 1993), both in field and in clusters. Although the only sample of ELGs photometrically analysed, as the UCM survey, in the *r* band, by Bothun et al. (1989) has not been morphologically described, these authors claim that many galaxies in the Wasilewski sample show exponential profiles, confirming that early types galaxies are not frequent in ELGs samples selected by objective-prism technique.

6.2. Hubble types for special classes of galaxies

In the UCM subsample of 12 Seyfert galaxies, the five Sy 1 present S0 types whereas the seven Sy 2 show later morphologies (one Sa, three Sb and three Sc). This result confirms the predominance of early types among the Sy 1 and that of later types for Sy 2, as mentioned in the previous section.

The global frequency of barred systems in the ELGs of the UCM survey is 9% (18/194). Considering the Hubble types, 5% of the barred UCM galaxies are S0, 17% are Sa, 50% are Sb and the remaining 28% are Sc⁺, showing then a predominance of later types, as in the whole UCM sample. And about the bar/ring-activity connection, in our subsample of 12 Seyfert galaxies, only two cases (11%) clearly present an underlying barred structure (UCM 1659+2928 and UCM 0142+2137) and only one (UCM 2329+2500) shows a ring (Gallego et al. 1994).

Since star formation is enhanced by interactions between galaxies, we have investigated the frequency of interacting systems in our survey of ELGs. In our exploration we have detected 11 systems with interacting features. These systems involve 16 UCM ELGs, since the remaining galaxies does not show obvious emission in our objective-prism plates due to either their weakness or their saturation on the plates. In the interacting UCM

subsample of ELGs, 25% are S0, 19% Sa, 44% Sb and the remaining 12% are Sc⁺. Therefore, this subsample is also dominated by late-type galaxies, although there is a relatively high fraction of S0's. This fact is in agreement with the idea proposed by some authors (e.g. van der Bergh 1990) in the sense that, at least in some cases, the S0 morphology could have been caused by present encounters or past mergers. About the presence of bars, two UCM interactive galaxies (13%), the SBb UCM 1537+2506N and the SBa UCM 1537+2506S, are barred. Concerning the occurrence of activity, 2 out of 16 interacting galaxies (13%) present Seyfert features (the Sy 1 UCM 1701+3131 and the Sy 2 UCM 1259+2934). Finally, in a global context, the 16 galaxies implicated in interactions represent 8% of the total UCM ELGs, a fraction lower than the 22% found in the Wasilewski sample (Bothun et al. 1989), although the problems in the characterization of an interaction only by means of the inspection of the image must be considered.

6.3. S0 galaxies in the UCM ELGs

Whereas in clusters, a fraction of S0 galaxies could have gas-depleted spirals as progenitors, the origin of field S0 galaxies is uncertain (gas loss by galactic winds, interactions, etc.). We have analysed the S0 nature from a photometric point of view. Following Kent (1985), if S0's were originated from spirals depleted of their gas contents, this would lead to a cessation of star formation and then a dimming in the disks would also be observed. Therefore, an increase in the B/D ratio should be expected, as well as a decrease in the effective radius a_e and the mean surface brightness inside the effective radius $\bar{\mu}_e$. To the contrary to this idea of S to S0 evolution, Kent (1985) finds that μ_0 values of S0's are similar to those of spirals, μ_e values for bulges of S0 are about 2 mag brighter than those of Sc galaxies and that S0's exhibit $\bar{\mu}_e$ not weaker but brighter than those of spirals.

Even taking into account that our S0 galaxies have been selected by the presence of emission features in their spectra and considering the problems in the $B + D$ decomposition for this morphological type, we have studied this problem in the way proposed by Kent (1985). For the subsample of 17 S0 field galaxies we have found μ_0 values about 1.5 mag weaker than the UCM sample mean, but me values for bulges of S0 about 2 mag brighter than the mean, in agreement with Kent (1985). Concerning the length-scales for S0, we find a 20% decrease in r_e and a 68% increase in d_L when comparing with the mean values of the whole sample. Globally combined these effects, they cause the expected rise in the B/D ratio when passing from S to S0 galaxies. We have also found that the absolute magnitudes for S0 bulges, M_{Bulge} , are 1.6 mag brighter than the mean, whereas for the disks, M_{Disk} , they are 0.3 mag weaker than the mean of the sample. Finally, concerning the mean photometric parameters, we find for

S0's a drop of 13% in a_e and a brightening of 0.8 mag in $\bar{\mu}_e$ when comparing with the mean values for the UCM sample. In summary, although some of these results (especially the rises in μ_e and $\bar{\mu}_e$) support the ideas of Boroson (1981) and Kent (1985) against a S to S0 evolution in the field, others like a dimming in μ_0 for S0 are opposed to the result found by Kent (1985).

Table 3. Morphological comparison for ELGs between UCM and Gisler samples

Sample	E-S0	Sa	Sb	\geq Sc
Gisler	12 %	10 %	28 %	50%
UCM	10 %	25 %	30 %	35%

6.4. Comparison with Gisler sample

We compare here the UCM distribution of morphologies with that from previous studies. Since a morphological classification of galaxies selected by an objective-prism technique has been never undertaken so far, we have only compared our results with a sample of emission-line galaxies, selected by different methods. This is the sample of Gisler (1978), with 1316 galaxies (799 with emission). For comparison purposes we have binned Gisler morphological groups in our types E-S0, Sa, Sb and Sc⁺ (Sc or later). Table 3 shows that whereas the fractions of early galaxies (E-S0) and Sb type are very similar in both samples, we have a higher detection of Sa and a lower fraction of late-type (Sc⁺) galaxies. Our excess in Sa detections could be due to the frequent presence of nuclear emission in our sample. And, concerning the excess of Sc⁺ galaxies in the Gisler sample, it could come from the inclusion by this author of a not specified number of galaxies with multiple or unclassified morphologies. These difficulties in the morphological comparisons reveal the necessity of Hubble classifications for recent objective-prism explorations of ELGs as Wasilewski or Michigan surveys.

7. Influence of the environment in the presence and nature of ELGs

Since one of the objective-prism plates in the UCM survey covers the Coma cluster, we have analyze whether there is some influence of the environment in the frequency and nature of ELGs when comparing our samples in the field and in the cluster medium. This comparison must be considered as indicative since only a cluster is concerned in our observations.

7.1. Frequency and Hubble types of UCM ELGs in the Coma cluster

A previous analysis for the UCM survey (Zamorano et al. 1996) has shown that the fraction of detected ELGs relative to the total number of Zwicky galaxies in the same area of the sky goes down from 15% to 7% when comparing the field to the Coma cluster. This decrease in the fraction of ELGs in a dense medium, has been previously observed e.g. by Gisler (1978), Dressler et al. (1985) and Salzer et al. (1988), and it is related to the morphology-density relation (Dressler 1980): in a denser environment, a lower quantity of gas available corresponds to a higher predominance of early-type (E-S0) galaxies and, therefore, to a reduced presence of star-forming systems, dominated by late-type galaxies.

In Table 4 we present the morphological distribution for UCM galaxies in the field (157) and in the Coma cluster (37). Although in both environments the sample is dominated by late-type galaxies, when passing from the field to the cluster we observe drastic reductions in the fraction of early-types (E-S0) ELGs and late-type spirals (Sc⁺), although there is an increase in the Irr class. On the other hand, the frequency of Sa remains almost the same, and there is a significative increase in the fraction of Sb galaxies. Since this analysis has never been undertaken in other surveys of ELGs selected by an objective-prism technique, we only have been able to make a comparison with some samples of galaxies selected by other criteria. So, the same effects, excluding the reduction in the late-type galaxies, were observed by Gisler (1978). However, Moss & Whittle (1993) found no difference in the fraction of early-types ELGs when comparing the field with a cluster environment, although they found the same effect in the reduction of late-type spirals.

7.2. Interpretation of the observed results in Coma

Although the distribution of morphological types in the Coma cluster has been investigated by several authors (e.g. Doi et al. 1995), the interpretation of our results is not straightforward since it should not be forgotten that our subsample in Coma does not correspond to a general sample of normal galaxies, but of galaxies selected by the presence of emission features in their objective-prism spectra. Thus, in our subsample of Coma galaxies, we are biased to star forming systems (i.e. late-type galaxies) that, as proposed by e.g. Bothun et al. (1984), could afterwards evolve to early-type galaxies due to a sweeping of their gas when falling to the cluster center. Concerning the excess in the Irr type in Coma, although Gallagher & Hunter (1989) propose a lower star formation efficiency in the cluster Irr compared to field Irr, we find that 4 out of 5 (80%) of the irregular galaxies found in the UCM survey belong to the Coma cluster. The high fractions of Sa-Sb types observed in our survey could be explained by the excess of

red color in the central parts of the disk galaxies observed by Gavazzi et al. (1990). Finally, concerning the presence of interactions, we find a reduction from 9% to 5% when comparing the fraction of galaxies with signs of interactions in the field to that in the Coma cluster. This could suggest that in the regions of low local galaxy density, the smaller relative velocity in the encounters between galaxies could favour the interacting systems. Unfortunately, since the question concerning whether the star formation rate is lower (Gisler 1978) or higher (Maia et al. 1994) in denser environments compared to the field has not been answered yet, and considering the instrumental effect of the different selection techniques (Moss & Whittle 1993), any comparison must be taken with caution.

8. Summary

We have carried out an extensive Thuan & Gunn *r* CCD imaging of the Universidad Complutense de Madrid (UCM) sample of emission-line galaxies (ELGs), which have been selected from the presence of H α emission in low resolution objective-prism plates. In Paper I (Vitores et al. 1996) we presented the observations and basic reductions for the set of 212 objects from the UCM survey. Also, a set of photometric parameters has been calculated to obtain, by using simultaneously five criteria, a morphological classification of the objects. In this paper we have analysed the statistical properties of the sample and their interpretations.

For the first time in a sample of ELGs, a complete analysis of the photometric data and morphological types has been carried out. For the whole UCM sample, we present histograms for all the computed parameters: magnitudes, sizes, *B+D* decomposition values, mean photometric radii and surface brightnesses, and concentration indices. Some correlations between these photometric parameters have also been analysed.

The statistical analysis of the obtained results has shown that, photometrically, our survey selects objects fainter (in apparent magnitudes) than those of color surveys and confirms that objective-prism surveys carried out in the blue and in the red bands are complementary to each other. In absolute values, our survey also confirms the advantage of emission-line surveys in detecting objects intrinsically fainter than those found from other selection techniques. The UCM survey attains median values for the whole sample of $M_r = -20.6$ and $D_{24} = 19.4$ kpc. Morphologically, the sample comprises a very heterogeneous population of galaxies, although it is dominated by late-type galaxies (66% Sb or later), as it is confirmed by means of several criteria. We propose a set of 9 candidates to BCD (blue compact dwarf) galaxies, five of them previously unknown. We have found that roughly 10% of the ELGs of the sample exhibit typical parameters of S0 types and, in the subsample of the 16 UCM interacting galaxies, 25% are classified as S0's.

Table 4. Morphological comparison for UCM galaxies in the field and in the Coma cluster. Between parenthesis is presented the total number of galaxies for each category

Subsample	E-S0	Sa	Sb	Sc ⁺	Irr	BCD
Field	11% (18)	24% (38)	28% (44)	31% (48)	1% (1)	5% (8)
Cluster	3% (1)	27% (10)	41% (15)	16% (6)	11% (4)	3% (1)

Finally, our study of the rich Coma cluster by means of the objective-prism technique allows the detection of a population of late Hubble types. The photometric analysis has found that 71% of the UCM ELGs in the Coma cluster are classified as Sb type or later.

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